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ASKAP: Setting the scene

Max Voronkov ASKAP Computing 23rd August 2010



ASKAP overview



- Located at radio-quiet site approx. 300 km inland from Geraldton
- Array of 36 12m antennas with phased array feeds (PAF)
- Initially 6 antennas: BETA (Boolardy Engineering Test Array)
- First antenna on site since last year, already used for some science (VLBI with other Australian antennas and Warkworth in New Zealand)



ASKAP 3-axis antenna mount

• 3-axis mount allows us to keep beam pattern fixed on the sky



ASKAP: General project news

- Antennas ahead of schedule
 - BETA antennas (1-6) to arrive to WA within a month
- Same hardware for beamformer, correlator and tide array unit
 - Digital team redesigned hardware for Virtex 6 FPGA
 - Virtex 7 could be a game changer (direct sampling + 4 boards instead of 32)
- Ten survey science projects
 - Two high priority projects (EMU, Wallaby)
 - Simulations to ensure software is ready
- PAF is the main technical risk
 - New technology, fundamentals to learn
 - Aggressive timescale
 - Economical production
 - Performance requirements
- Scaling is another risk





Calibration & Imaging challenges

- Strong sources contaminating the data through primary beam sidelobes
 - We have 3-axis mount which keeps sidelobes fixed
 - Beam variations due to PAF instabilities could be a problem
- Wide field calibration
 - Ionosphere is benign at frequencies about 1 GHz
 - PAF is stabilized in hardware (noise sources)
 - Software calibration is per synthetic beam
- Wide field imaging
 - Take direction-dependent effects via convolution functions
- Wide field deconvolution
 - Subtraction of the local sky model from uv-data
 - Joint processing of the full field of view
 - S/N-based cleaning (eventually MSMF algorithm)



Calibration & Imaging challenges - 2

- Mosaicing in full polarization
 - Polarisation properties of each beam will be taken care of by adding an extra dimension to convolution functions
- Mosaicing with different primary beams
 - Comes out naturally in our approach to mosaicing (we planned for this up front designing our software, but haven't tried this case yet in practice)
- Large data volumes (LDV) pipeline processing
 - Central Processor of ASKAP will reduce data on-the-fly, astronomers are not expected to touch uv-data
- Large data volumes (LDV) data formats
 - At this stage we use Measurement Sets. In any case, we plan to write a tool exporting the data into MS to assist with debugging (e.g. using casa)



Calibration & Imaging challenges - 3

- Large data volumes (LDV) processing power limitations and shortcuts (e.g. algorithm and data compression) needed
 - Shortcuts to ensure single iteration over data is sufficient
 - Replaced traditional weighting schemes with post-gridding preconditioning (e.g. Wiener filter)
 - Assumed a good instantaneous uv-coverage
- Sky models: greater sophistication in specification
 - Plan to reuse LOFAR approach
 - Not much research done so far
- Solvability (cal): enough calibrators?
 - ASKAP field of view has on average 56 Jy of flux
 - With the target performance figures / 5 min solution interval, it allows to calibrate gain amplitudes with the 2% accuracy and phases with a few degrees accuracy

• The impact on the dynamic range is not clear

Calibration & Imaging challenges - 4

- Time and frequency dependence of calibration parameters
 - Predict forward approach (no interpolation)
 - Frequency dependence: bandpass, leakages per coarse
 - (1 MHz) channel
- Full pol imaging
 - Specify polarisation of the primary beams via convolution functions (extra dimansions)
- On-the-fly mapping (ask Gerry what this means)- Long baselines / large fields of view: dumping fast enough
 - ASKAP will deliver eventually a continuum map every 5 seconds (for transient search)





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Thank you

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