

The Cosmic Lens All-Sky Survey parent population – I. Sample selection and number counts

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ABSTRACT

We present the selection of the Jodrell Bank Flat-spectrum (JBF) radio source sample, which is designed to reduce the uncertainties in the Cosmic Lens All-Sky Survey (CLASS) gravitational lensing statistics arising from the lack of knowledge about the parent population luminosity function. From observations at 4.86 GHz with the Very Large Array, we have selected a sample of 117 flat-spectrum radio sources with flux densities greater than 5 mJy. These sources were selected in a similar manner to the CLASS complete sample and are therefore representative of the parent population at low flux densities. The vast majority (~ 90 per cent) of the JBF sample are found to be compact on the arcsec scales probed here and show little evidence of any extended radio jet emission. Using the JBF and CLASS complete samples, we find the differential number counts slope of the parent population above and below the CLASS 30-mJy flux density limit to be -2.07 ± 0.02 and -1.96 ± 0.12 , respectively.

Key words: surveys – galaxies: active – radio continuum: galaxies.

1 INTRODUCTION

Gravitational lensing statistics are a useful tool for constraining the cosmological parameters, or alternatively, investigating the global properties of lensing galaxy populations (Turner, Ostriker & Gott 1984; Carroll, Press & Turner 1992; Fukugita et al. 1992; Kochanek 1996a; Helbig et al. 1999; Chae et al. 2002; Chae 2003, 2005, 2007; Chae & Mao 2003; Davis, Huterer & Krauss 2003; Kuhlen, Keeton & Madau 2004; Mitchell et al. 2005; Chae, Mao & Kang 2006). However, both applications require a complete sample of gravitational lenses, drawn from a parent population¹ with a well-defined selection function (Kochanek 1996b). Such samples can be best obtained at radio wavelengths since dust obscuration within the lensing galaxy, which plagues optical-based surveys, is not a factor and the high angular resolutions available can detect subarcsec image separations. Radio surveys can also be complete, flux density limited and carried out efficiently, with almost uniform sensitivity and resolution. Moreover, radio surveys with a flat-spectrum radio source parent population have been successful because the simple compact structure of the background source allowed the straightforward identification of galaxy-scale image splitting (typically ~ 1 arcsec) with high sensitivity instruments such as the Very Large Array (VLA).

The Cosmic Lens All-Sky Survey² (CLASS; $S_{4.85} \geq 30$ mJy; Browne et al. 2003; Myers et al. 2003) forms the largest, statistically complete sample of radio-loud gravitational lens systems currently available. A complete sample of 11 685 flat-spectrum radio sources (the exact selection criteria for this parent population sample are given in Section 2) was observed with the VLA at 8.46 GHz in A configuration (resolution of ~ 0.2 arcsec). Those sources which were found to have multiple components with Gaussian full width at half-maximum (FWHM) ≤ 170 mas, flux density ratios $\leq 10:1$ and separated by ≥ 300 mas in the CLASS 8.46-GHz VLA images were followed-up as potential gravitational lensing candidates. Further observations with optical telescopes and high-resolution radio arrays confirmed the lensing hypothesis for 22 gravitational lens systems during the course of CLASS. Of these systems, 13 form a well-defined statistical sample of gravitational lenses from a parent population of 8958 flat-spectrum radio sources. This results in a CLASS lensing rate of 1:689. Further details of the CLASS gravitational lens systems, and the procedures used to discover them, can be found in Browne et al. (2003).

A thorough analysis of the CLASS gravitational lensing statistics found, for a flat-Universe with a classical cosmological constant ($w = -1$), $\Omega_{\Lambda} = 0.69^{+0.14}_{-0.27}$ at the 68 per cent confidence level (Chae et al. 2002; Chae 2003). This result, which is consistent

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¹ Throughout, we define the parent population as a sample of objects which could potentially be gravitationally lensed. The selection of such a sample is independent of the lensing probability of each individual object.

² The Jodrell Bank–VLA Astrometric Survey (JVAS; $S_{4.85} \geq 200$ mJy; Patnaik et al. 1992; Browne et al. 1998; Wilkinson et al. 1998; King et al. 1999) forms a brighter subsample of CLASS.

with the findings from SN1a (e.g. Riess et al. 2004), large-scale structure (e.g. Cole et al. 2005) and cosmic microwave background (e.g. Spergel et al. 2006) data, provides further independent evidence for the cosmological concordance model. Furthermore, the CLASS gravitational lensing statistics have also been used to investigate the global properties of the lensing galaxy population. Chae et al. found the characteristic velocity dispersion for the early- and late-type galaxy populations to be $\sigma_*^{(e)} = 198_{-37}^{+58}$ km s⁻¹ and $\sigma_*^{(l)} = 117_{-31}^{+45}$ km s⁻¹ at the 95 per cent confidence level (see also Chae 2003; Davis et al. 2003). The projected mean ellipticity for the early-type population, based on the relative numbers of quadruple and doubly imaged CLASS gravitational lens systems, was found to be $\bar{f} < 0.83$.

The analyses described above required the number density of the parent population as a function of flux density to be established. This is because the derived constraints on $\Omega_M - \Omega_\Lambda$ depend on a knowledge of the lensing optical depth as a function of the background source redshift (e.g. Turner et al. 1984). Unfortunately, the flat-spectrum radio source luminosity function was not well known, and measuring the redshifts of the 11 685 sources in the CLASS complete sample was not practical. Therefore, subsamples of flat-spectrum radio sources, selected in a similar manner to the CLASS complete sample, were formed within progressively lower flux density bins. At high flux densities, the parent population redshift information was taken from the Caltech–Jodrell Bank Flat-spectrum survey (CJF; $S_{4.85} \geq 350$ mJy; Taylor et al. 1996). The complete CJF sample consists of 293 flat-spectrum radio sources, for which, 261 redshifts have been obtained (Vermeulen & Taylor 1995; Vermeulen et al. 1996; Henstock et al. 1997, also unpublished). A redshift survey of 69 sources from the JVAS sample by Falco, Kochanek & Muñoz (1998) has provided 55 redshifts in the intermediate flux density range, 200–250 mJy at 4.85 GHz (see also Muñoz et al. 2003). Redshift information for the parent population at the CLASS flux density limit was reported by Marlow et al. (2000), who measured 27 redshifts from a sample of 42 flat-spectrum radio sources with 4.85-GHz flux densities between 25 and 50 mJy. The mean redshift of each of these flat-spectrum radio source samples is $\bar{z} \sim 1.25$, suggesting little change in the mean redshift with flux density.

However, since gravitational lensing increases the apparent flux density of the background source, many lensed sources will come from a population of radio sources with flux densities below the CLASS flux density limit. Therefore, our knowledge of the flat-spectrum radio source luminosity function must be extended below 25 mJy to a few mJy (based on the source magnifications calculated from lens galaxy mass modelling). We have therefore undertaken a study of the flat-spectrum radio source population at the mJy level; hereafter referred to as the Jodrell Bank Flat-spectrum (JBF) radio source survey. The aim of this study is to reduce the uncertainties in the CLASS gravitational lensing statistics arising from the parent population luminosity function.

Since this project began, Muñoz et al. (2003) have extended their work on the redshift distribution of flat-spectrum radio sources down to ~ 3 mJy. They find the mean redshift of their sample of 33 flat-spectrum radio sources with ~ 5 GHz flux densities between 3 and 20 mJy to be $\bar{z} \sim 0.75$ (42 per cent completeness). This mean redshift is significantly lower than the trend reported from the subsamples of flat-spectrum radio sources selected from the CJF, JVAS and CLASS surveys. The implications of such a low mean redshift for the parent population at low flux densities on the CLASS lensing statistics is to push Ω_Λ to ~ 1 for a flat-Universe, which is inconsistent with the concordance model. In a companion paper (McKean

et al., in preparation), we will present the optical and near-infrared followup of a small subsample of JBF sources which will show that the mean redshift of the parent population is nearer $\bar{z} \sim 1.2$ at low flux densities. The focus of this paper, which is the first in a series of papers investigating the flat-spectrum radio source population at the mJy level, is to present the selection of the JBF sample and the number counts of the CLASS parent population.

In Section 2, we review the strict selection criteria of the CLASS complete and statistical samples. New 4.86- and 8.46-GHz observations from a VLA pseudo-survey that were used to select the JBF sample are presented in Section 3. In Section 4, we discuss the radio morphologies of the 117 flat-spectrum radio sources in the JBF sample. We also present our analysis of the CLASS parent population differential number counts and discuss the implications for the CLASS gravitational lensing statistics in Section 4. We end with a summary of our findings in Section 5.

2 THE CLASS COMPLETE AND STATISTICAL SAMPLES

To be truly representative of the CLASS parent population, the JBF sample needed to be selected in an identical manner to the flat-spectrum radio sources observed by CLASS. Therefore, we first present a brief review of the selection criteria for the CLASS complete and statistical samples before discussing the selection of the JBF sample.

The well-defined CLASS complete sample was selected using the 1.4-GHz NVSS (National Radio Astronomy Observatory Very Large Array Sky Survey; Condon et al. 1998) and the 4.85-GHz GB6 (Green Bank 6 cm; Gregory et al. 1996) catalogues to find all flat-spectrum radio sources with

- (i) $\alpha_{1.4}^{4.85} \geq -0.5$ where $S_\nu \propto \nu^\alpha$,
- (ii) $S_{4.85} \geq 30$ mJy,
- (iii) $0^\circ \leq \delta \leq 75^\circ$ and
- (iv) $|b| \geq 10^\circ$.

The CLASS complete sample was selected by finding all sources with $S_{4.85} \geq 30$ mJy from the GB6 catalogue in the area of sky defined above. These sources were then cross-correlated with the NVSS catalogue (CATALOG39). All 1.4-GHz flux density within 70 arcsec of the GB6 position was summed and used to determine the two-point spectral index of each source. There are 11 685 flat-spectrum radio sources in the CLASS complete sample within a sky region of 4.96 sr. This sample was then observed with the VLA in A configuration at 8.46 GHz during CLASS. Those sources which were found to have a total 8.46-GHz flux density of $S_{8.46} \geq 20$ mJy formed the CLASS statistical sample. The 20-mJy cut-off was applied to ensure that all sources with multiple components and flux density ratios less than 10:1 would be detected by the VLA. There are 8958 sources in the CLASS statistical sample. The difference in the number of sources in the complete and statistical samples is mainly due to the 20-mJy cut-off (2418 sources). Bandwidth smearing (217 sources), extended sources (81 sources) and failed observations (11 sources) account for the remainder. A full discussion of the selection of the CLASS complete and statistical samples, and the subsequent CLASS VLA 8.46-GHz observations can be found in Myers et al. (2003).

3 SAMPLE SELECTION

Due to the magnification of the background source by gravitational lensing, we needed to determine the number counts and redshift

distribution of the parent population below the CLASS 30-mJy flux density limit at 4.85 GHz. Therefore, we selected a representative sample of faint flat-spectrum radio sources which is complete to 5 mJy. We now discuss the selection of the JBF sample.

3.1 The NVSS selected sample

GB6 could not be used as the primary source catalogue because the JBF sample would include flat-spectrum radio sources with ~ 5 GHz flux densities down to 5 mJy (recall that the GB6 catalogue is flux density limited to $S_{4.85} \geq 18$ mJy). Ideally, we would carry out our own, deeper sky survey at ~ 5 GHz to identify a flux density limited sample of faint flat-spectrum radio sources. However, this process would be observationally expensive. Therefore, using the VLA at 4.86 GHz, we undertook a targeted pseudo-survey of a well-defined sample of radio sources selected from the NVSS catalogue ($S_{1.4} \geq 2.5$ mJy) within a restricted region of the sky. From these 4.86-GHz pseudo-survey observations, we established a sample of NVSS-selected radio sources which met the CLASS two-point spectral index criteria ($\alpha_{1.4}^{4.86} \geq -0.5$) and had $S_{4.86} \geq 5$ mJy. This process is slightly different to the one used for the selection of the CLASS complete sample (see Section 2). Therefore, we now discuss any possible bias which the 4.86-GHz pseudo-survey may have introduced.

The NVSS $S_{1.4} \geq 2.5$ mJy limit was chosen to ensure that a sample of flat-spectrum radio sources with $S_{4.86} \geq 5$ mJy was selected. However, this limit also imposed on the pseudo-survey a bias against faint and highly inverted flat-spectrum radio sources with $\alpha_{1.4}^{4.86} \geq 0.56$ (e.g. for a 5-mJy source at 4.86 GHz). Assuming that the spectral index distribution of the flat-spectrum radio sources found by the pseudo-survey is the same as for the CLASS complete sample (see fig. 3 in Myers et al. 2003), we would expect 9.4 per cent of the sources to have $\alpha_{1.4}^{4.86} \geq 0.5$. This does not mean that the pseudo-survey would not detect any of these inverted radio sources; as we will see in Section 4.1, 6 per cent have $\alpha_{1.4}^{4.86} \geq 0.5$. It is only the few highly inverted radio sources (3.4 per cent) at the 5-mJy limit of the pseudo-survey which would be missed.

The GB6 survey was conducted with the old 300-ft (91-m) Telescope at Green Bank which had a beam size of ~ 3.5 arcmin, whereas our 4.86-GHz pseudo-survey observations were carried out using the VLA, with a beam size of only a few tens of arcsec. This change in resolution will result in two effects. First, the increase in resolution introduced the possibility of the pseudo-survey observations resolving several discrete sources that would have been identified as a single radio source by GB6. When this occurred, we summed the 4.86-GHz flux density of the separate sources to make a single ‘radio’ source (the details of this process are given in Section 3.3). Secondly, the higher resolution provided by our interferometric VLA observations could result in extended radio emission being partially or completely resolved out. However, since the aim of this project is to select a sample of flat-spectrum radio sources, which are typically

compact, we expect this to have a negligible effect on our sample completeness.

The number of NVSS radio sources with $S_{1.4} \geq 2.5$ mJy is approximately 44 sources deg^{-2} . Therefore, to define a complete low flux density sample which was also straightforward to follow up at optical wavelengths, sources were selected from 16 circular fields with radii ranging from 0.3° to 1° within the region of sky $13^{\text{h}} \lesssim \alpha \lesssim 8^{\text{h}}$ and $\delta \sim 0^\circ$. Where possible, fields were chosen to coincide with the Anglo-Australian Observatory 2dF Galaxy Redshift Survey (Folkes et al. 1999) in the hope that some of the flat-spectrum radio sources would have measured redshifts. There are 1299 sources in the complete $S_{1.4} \geq 2.5$ mJy sample within a sky area of 29.3 deg^2 ($\equiv 8.93 \times 10^{-3} \text{ sr}$).

3.2 VLA 4.86-GHz pseudo-survey observations

The complete NVSS selected $S_{1.4} \geq 2.5$ mJy sample was observed at 4.86 GHz with the VLA in CnD configuration on 1999 March 02 (6 h) and 1999 March 05 (4 and 3 h), and in D configuration on 1999 May 21 (12 h). Each source was observed for 45 or 50 s, using a 10-s correlator integration time. The data were taken through two 50-MHz IFs, which were centred at 4.835 and 4.885 GHz, respectively. 3C 286 and 3C 48 were used as the primary flux density calibrators and suitable phase reference calibrators, selected from the JVAS catalogue, were observed every 15–30 min. The typical beam size was $\sim 20 \times 13 \text{ arcsec}^2$ with an rms map noise $\sim 300 \mu\text{Jy beam}^{-1}$. A summary of the VLA 4.86-GHz pseudo-survey observations is given in Table 1.

The data were calibrated and edited in the standard way using the AIPS (Astronomical Image Processing Software) package. To ensure that the imaging of the data was carried out in an efficient and consistent manner, all 1299 pointings were mapped within the Caltech Very Long Baseline Interferometry difference mapping package (DIFMAP; Shepherd 1997) using a modified version of the CLASS mapping script (Myers et al. 2003). The script automatically detected and *cleaned* surface brightness peaks above $1.5 \text{ mJy beam}^{-1}$ which had a signal-to-noise ratio greater than 6 (typically $\gtrsim 2.4 \text{ mJy beam}^{-1}$), within a sky region of $2048 \times 2048 \text{ arcsec}^2$ in size around the phase centre. Natural weighting was used throughout to maximize the overall signal-to-noise ratio, and elliptical Gaussian model components were fitted to the data.

3.3 The JBF sample

The pseudo-survey observations were carried out to emulate what was done for the GB6 survey using the old 300-ft (91-m) Telescope at Green Bank. However, the GB6 survey has a beam size of ~ 3.5 arcmin, which is significantly larger than the pseudo-survey $20 \times 13 \text{ arcsec}^2$ beam size. This introduced the possibility of the pseudo-survey observations resolving discrete sources that would

Table 1. A summary of the number of sources observed, detected and found to have flat-spectra from the VLA pseudo-survey. The number of sources observed relates to the actual number of VLA pointings. Those sources which were within 70 arcsec of the VLA pointing and have a flux density of $S_{4.86} \geq 5$ mJy were classed as detections.

Date	Array	Integration time (s)	Sources observed	Sources detected	Percentage detected	Flat-spectrum sources	Percentage flat-spectrum
1999 March 02	CnD	50	333	112	34	34	30
1999 March 05	CnD	45	353	107	30	34	32
1999 May 21	D	45	613	199	32	49	25
Total			1299	418	32	117	28

have otherwise been identified as a single radio source by GB6. This issue was also confronted during the selection of the CLASS complete sample where the NVSS beam size (45 arcsec) was ~ 4 times smaller than the GB6 beam size. To overcome this relative beam size problem, Myers et al. (2003) added all the NVSS 1.4-GHz flux density within 70 arcsec of the GB6 position to determine the 1.4-GHz flux density of each ‘source’. We have adopted the same strategy for the pseudo-survey. The 4.86-GHz radio emission from those pseudo-survey sources within 70 arcsec of each other were added together to make a single radio source and entered into the 4.86-GHz pseudo-survey catalogue. As the pointings for the 4.86-GHz pseudo-survey observations were taken from the NVSS catalogue, there was the possibility that a source was detected in more than one field. When this occurred, the data from the nearest pointing were used. The 4.86-GHz catalogue was then cross-referenced with the NVSS catalogue. As with the selection of the CLASS complete sample, the total 1.4-GHz flux density within 70 arcsec of the 4.86-GHz position was added and used to determine the two-point spectral index of each source.

The pseudo-survey catalogue contains 736 sources detected at 4.86 GHz with the VLA. Of these sources, 418 are in the flux density limited sample of $S_{4.86} \geq 5$ mJy. This results in a source density above 5 mJy of about 14 sources deg^{-2} . For the pseudo-survey, this equates to one source every 10^3 – 10^4 beam areas. For a radio source population whose differential number counts are described by a power law with an index of 2 (see Section 4.2), we would expect confusing sources (i.e. those at a density of 1 per 20 beam areas) to contribute about 0.1 mJy to the flux density of a 5-mJy source. This is well within the observational uncertainties of the pseudo-survey flux densities. Therefore, source confusion will have a negligible effect on the pseudo-survey catalogue at the 5-mJy flux density limit. The total number of flat-spectrum radio sources defined by the CLASS two-point spectral index criteria within the $S_{4.86} \geq 5$ mJy flux density limited sample is 117 sources. It is these 117 flat-spectrum radio sources which form the JBF sample. A summary of the number of sources observed, detected and found to have flat radio spectra during each VLA observing run is given in Table 1. We find no significant differences in the results from the three observing periods. The positions, flux densities and spectral indices of each flat-spectrum radio source in the JBF sample are given in Table 2.

3.4 VLA 8.46-GHz observations

The final step of the JBF sample selection process was the application of the same observational biases and filters imposed on the CLASS statistical sample. This was done by observing the JBF sample with the VLA at 8.46 GHz in A configuration on 1999 June 29. The higher resolution 8.46-GHz observations also provided the accurate positional information required for future optical and infrared followup work, and determined if there were any gravitational lensing candidates in the JBF sample. However, only 59 JBF sources were observed because of an error in an initial reduction of the 4.86-GHz pseudo-survey data prior to the 8.46-GHz observations. Each source was observed for 100 s. A 10-s correlator integration time was used through two 50-MHz IFs, which were set to 8.435 and 8.485 GHz, respectively. As before, 3C 286 was used as the primary flux density calibrator and phase referencing was carried out with suitable JVAS sources. The typical beam size was $\sim 0.7 \times 0.2$ arcsec² and the rms map noise was ~ 180 $\mu\text{Jy beam}^{-1}$. The data were reduced using AIPS. Mapping and self-calibration were carried out within DIFMAP. Natural weighting was used, and elliptical Gaussian model components were fitted to the data.

All 59 sources were detected and have compact structures (Gaussian FWHM ≤ 170 mas). The positions, flux densities and spectral indices for each source are given in Table 3. Only one source was found to have multiple components. JBF.041 has two compact components (Gaussian FWHM of 60 and 120 mas) separated by 1.47 arcsec. Independently of this work, JBF.041 was identified as a gravitational lens candidate from the PMN survey (Parkes–Massachusetts Institute of Technology–National Radio Astronomy Observatory; Griffith & Wright 1993). Extensive radio and optical observations by Winn et al. (2002) have shown PMN J1632–0033 (JBF.041) to be a gravitational lens system, with three lensed images of a $z = 3.42$ quasar (see also Winn, Rusin & Kochanek 2003, 2004).

4 DISCUSSION

4.1 Radio morphologies and extended emission

We have investigated the morphological properties of the JBF sample by classifying each flat-spectrum radio source as either unresolved (U) or extended (E) in Table 2. Unresolved radio sources are those consisting of a single radio component (within a 70-arcsec search radius from the brightness peak) with a model Gaussian FWHM which is smaller than the observed beam size of the VLA (also given in Table 2). The remainder are considered extended.

Our analysis of the 4.86-GHz VLA model-fitting data finds 85 per cent of the JBF sample to have unresolved structures. Evidence for extended emission is found in 15 per cent of the radio sources. The large fraction of unresolved point sources in the JBF sample is not unexpected – the high selection frequency, coupled with the tight constraint on the source spectral index should have produced a sample of core-dominated radio sources. In Fig. 1, we show the spectral index distribution of the complete JBF sample (solid line). The $\alpha_{1.4}^{4.86} \geq -0.5$ spectral index cut can be clearly seen in the distribution. Of the full JBF sample, 32 per cent have a rising radio spectrum between 1.4 and 4.86 GHz (i.e. $\alpha_{1.4}^{4.86} \geq 0$) and only 6 per cent are highly inverted (i.e. $\alpha_{1.4}^{4.86} \geq 0.5$). The total sample of 117 flat-spectrum radio sources has a mean spectral index of -0.09 with an rms of 0.31 and a median spectral index of -0.15 . We also show in Fig. 1, with the broken line, the spectral index distribution of those sources which are considered extended. The broken line effectively divides each spectral index bin into the contribution from unresolved and extended radio sources. It is apparent that the extended radio sources tend to have on average steeper radio spectra (mean spectral index is -0.22 with an rms of 0.25; median spectral index is -0.23) when compared to the unresolved population (mean spectral index is -0.07 with an rms of 0.32; median spectral index is -0.10). The steeper spectra are likely caused by the presence of jet emission in the extended sources, or due to contamination from another independent (steep spectrum) radio source within 70 arcsec of the brightness peak.

We have searched for any evidence of extended jet emission in the JBF sample by inspecting the radio maps of those sources observed during the course of the 1.4-GHz FIRST survey (Faint Images of the Radio Sky at Twenty centimetres; Becker, White & Helfand 1995; beam size ~ 5 arcsec). We found that only 33 of the 117 JBF sources have FIRST radio maps available due to the limited sky coverage of the FIRST survey. The mean spectral index of these 33 JBF sources is -0.11 , with 18 per cent (six sources) defined as extended in Table 2. Therefore, the 33 sources appear to form a representative subsample of the JBF catalogue (cf. with the mean spectral index and

Table 2. The JBF 4.86-GHz catalogue. The survey name of each flat-spectrum radio source is given in Column 1. The J2000 right ascension and declination are listed in Columns 2 and 3, respectively. For each source, the peak surface brightness (Column 4) and the integrated flux density (Column 5) from model fitting to the uv data is reported. The radio morphology of each JBF source has been classified as either unresolved (U) or extended (E) in Column 6. The particulars of the 4.86-GHz observation of each object are given in Columns 7–10. The 1.4-GHz NVSS flux density within 70 arcsec of the JBF position (Column 11) has been used to calculate the 1.4–4.86 GHz spectral index of each source in Column 12.

JBF name	RA ($^{\circ}$ $^{\prime}$ $^{\prime\prime}$)	Dec. ($^{\circ}$ $^{\prime}$ $^{\prime\prime}$)	I_{peak} (mJy beam $^{-1}$)	S_{int} (mJy)	Morph.	Observation date	Array	Beam size (arcsec 2 , PA)	σ_{map} (μ Jy beam $^{-1}$)	$S_{1.4}$ (mJy)	$\alpha_{1.4}^{4.86}$
JBF001	03 57 51.5324	+00 30 47.482	7.8 \pm 0.5	7.9 \pm 0.5	U	1999 May 21	D	20.7 \times 12.8, $-4^{\circ}7$	360	6.8 \pm 0.5	+0.12 \pm 0.08
JBF002	03 58 18.0268	+00 28 00.950	8.3 \pm 0.5	8.6 \pm 0.5	U	1999 May 21	D	20.7 \times 12.8, $-4^{\circ}4$	322	3.9 \pm 0.7	+0.64 \pm 0.15
JBF003	03 59 06.8984	-00 06 18.334	8.1 \pm 0.5	8.4 \pm 0.5	U	1999 May 21	D	20.8 \times 12.8, $-3^{\circ}7$	306	4.9 \pm 0.5	+0.43 \pm 0.09
JBF004	04 02 19.3023	-00 18 00.330	6.0 \pm 0.4	9.8 \pm 0.6	E	1999 May 21	D	20.8 \times 12.9, $+1^{\circ}6$	302	15.8 \pm 1.3	-0.38 \pm 0.08
JBF005	04 02 35.9708	+00 12 41.052	5.5 \pm 0.4	6.1 \pm 0.4	U	1999 May 21	D	20.7 \times 12.9, $+1^{\circ}7$	312	9.7 \pm 0.5	-0.37 \pm 0.07
JBF006	04 02 39.7719	+00 09 10.664	8.3 \pm 0.5	9.3 \pm 0.6	U	1999 May 21	D	20.7 \times 12.9, $+1^{\circ}8$	302	16.8 \pm 0.7	-0.45 \pm 0.06
JBF007	07 57 26.1501	+00 09 00.215	8.8 \pm 0.6	9.6 \pm 0.6	U	1999 May 21	D	23.8 \times 14.2, $-34^{\circ}0$	401	17.2 \pm 1.1	-0.49 \pm 0.07
JBF008	07 59 04.6890	+00 22 33.318	19.6 \pm 1.0	20.9 \pm 1.1	U	1999 May 21	D	24.1 \times 15.0, $-28^{\circ}6$	310	32.6 \pm 1.1	-0.36 \pm 0.05
JBF009	07 59 20.5743	-00 14 01.201	40.2 \pm 2.0	41.1 \pm 2.1	U	1999 May 21	D	24.1 \times 14.9, $-27^{\circ}6$	342	60.0 \pm 1.8	-0.30 \pm 0.05
JBF010	07 59 48.7479	-00 21 40.913	5.3 \pm 0.4	5.3 \pm 0.4	U	1999 May 21	D	23.6 \times 14.6, $-25^{\circ}1$	295	6.2 \pm 0.5	-0.13 \pm 0.09
JBF011	07 59 54.0151	+00 05 09.272	15.1 \pm 0.8	14.9 \pm 0.8	U	1999 May 21	D	23.2 \times 14.6, $-24^{\circ}4$	305	19.1 \pm 0.7	-0.20 \pm 0.05
JBF012	08 01 18.9023	+00 20 13.654	6.1 \pm 0.4	6.4 \pm 0.4	U	1999 May 21	D	22.7 \times 14.3, $-21^{\circ}6$	297	5.4 \pm 0.5	+0.14 \pm 0.09
JBF013	13 14 10.0897	+00 15 44.461	5.4 \pm 0.4	5.8 \pm 0.4	U	1999 March 05	CnD	33.1 \times 13.4, $-47^{\circ}9$	295	3.9 \pm 0.5	+0.32 \pm 0.12
JBF014	13 14 33.0293	+00 06 36.779	8.6 \pm 0.5	8.6 \pm 0.5	U	1999 March 05	CnD	31.9 \times 13.6, $-48^{\circ}9$	310	6.3 \pm 0.5	+0.25 \pm 0.08
JBF015	13 15 22.7524	+00 09 47.527	5.1 \pm 0.4	5.6 \pm 0.4	U	1999 March 05	CnD	31.0 \times 13.6, $-48^{\circ}5$	290	10.2 \pm 0.5	-0.48 \pm 0.07
JBF016	13 30 20.1927	+00 17 18.695	12.7 \pm 0.7	12.9 \pm 0.7	U	1999 March 02	CnD	20.5 \times 6.0, $-65^{\circ}9$	307	14.4 \pm 0.6	-0.09 \pm 0.05
JBF017	13 31 00.7274	-00 14 38.466	5.3 \pm 0.4	5.4 \pm 0.4	U	1999 March 02	CnD	19.5 \times 6.1, $-66^{\circ}8$	316	7.8 \pm 0.5	-0.30 \pm 0.08
JBF018	13 31 05.7737	-00 22 21.760	7.2 \pm 0.5	7.6 \pm 0.5	U	1999 March 02	CnD	19.1 \times 6.1, $-67^{\circ}2$	286	9.4 \pm 0.5	-0.17 \pm 0.07
JBF019	14 14 16.6021	+00 24 06.102	17.0 \pm 0.9	17.4 \pm 0.9	U	1999 March 05	CnD	41.1 \times 13.4, $-51^{\circ}8$	329	19.3 \pm 1.0	-0.08 \pm 0.06
JBF020	14 15 38.9146	-00 27 39.633	5.2 \pm 0.4	11.6 \pm 0.7	E	1999 March 05	CnD	36.7 \times 13.4, $-50^{\circ}4$	333	18.8 \pm 1.0	-0.39 \pm 0.06
JBF021	14 15 59.5661	+00 13 57.739	5.9 \pm 0.5	6.6 \pm 0.5	U	1999 March 05	CnD	35.0 \times 13.4, $-50^{\circ}3$	345	11.4 \pm 0.5	-0.44 \pm 0.07
JBF022	14 16 35.1606	+00 11 56.244	10.4 \pm 0.6	11.2 \pm 0.7	U	1999 March 05	CnD	33.1 \times 13.3, $-49^{\circ}7$	347	17.7 \pm 0.7	-0.37 \pm 0.06
JBF023	14 27 53.8192	+00 00 38.818	16.3 \pm 0.9	16.7 \pm 0.9	U	1999 March 02	CnD	23.3 \times 5.6, $-63^{\circ}0$	322	23.1 \pm 0.8	-0.26 \pm 0.05
JBF024	14 29 07.1693	+00 15 49.184	4.6 \pm 0.4	5.2 \pm 0.4	U	1999 March 02	CnD	21.7 \times 5.7, $-64^{\circ}1$	294	5.6 \pm 0.5	-0.06 \pm 0.09
JBF025	14 29 15.1193	-00 31 01.473	4.4 \pm 0.4	9.9 \pm 0.6	E	1999 March 02	CnD	21.6 \times 5.7, $-64^{\circ}0$	348	17.6 \pm 1.3	-0.46 \pm 0.08
JBF026	14 29 17.8883	-00 24 40.131	4.4 \pm 0.4	6.2 \pm 0.5	E	1999 March 02	CnD	21.5 \times 5.7, $-64^{\circ}1$	340	7.9 \pm 0.5	-0.19 \pm 0.08
JBF027	14 29 37.1268	-00 05 07.701	7.2 \pm 0.5	8.1 \pm 0.5	U	1999 March 02	CnD	21.3 \times 5.7, $-64^{\circ}3$	325	8.6 \pm 0.5	-0.05 \pm 0.07
JBF028	14 29 56.5964	-00 17 55.778	21.1 \pm 1.1	21.4 \pm 1.1	U	1999 March 02	CnD	20.7 \times 5.7, $-64^{\circ}7$	309	14.7 \pm 0.6	+0.30 \pm 0.05
JBF029	14 30 31.2568	-00 09 06.939	54.3 \pm 2.7	54.7 \pm 2.8	U	1999 March 02	CnD	20.3 \times 5.9, $-65^{\circ}5$	366	58.5 \pm 1.5	-0.05 \pm 0.05
JBF030	15 00 19.0920	+00 02 47.510	5.2 \pm 0.4	5.2 \pm 0.4	U	1999 March 05	CnD	33.9 \times 13.4, $-49^{\circ}9$	325	4.6 \pm 0.4	+0.10 \pm 0.09
JBF031	15 00 27.8126	+00 02 24.365	20.3 \pm 1.1	22.9 \pm 1.2	E	1999 March 05	CnD	33.7 \times 13.4, $-49^{\circ}8$	322	21.2 \pm 1.1	+0.06 \pm 0.06
JBF032	16 29 21.8169	+00 05 08.171	6.1 \pm 0.6	6.0 \pm 0.6	U	1999 March 05	CnD	77.2 \times 12.5, $-54^{\circ}4$	482	4.7 \pm 0.4	+0.20 \pm 0.11
JBF033	16 29 56.8045	+01 01 41.007	93.7 \pm 4.7	93.3 \pm 4.7	U	1999 March 05	CnD	59.3 \times 13.2, $-53^{\circ}9$	504	44.7 \pm 1.4	+0.59 \pm 0.05
JBF034	16 30 18.7229	-00 27 53.143	4.5 \pm 0.5	5.2 \pm 0.5	U	1999 March 05	CnD	54.3 \times 12.5, $-53^{\circ}2$	446	8.2 \pm 0.5	-0.37 \pm 0.09
JBF035	16 30 40.8391	+00 22 08.522	7.0 \pm 0.5	7.2 \pm 0.6	U	1999 March 05	CnD	49.5 \times 13.1, $-52^{\circ}9$	423	9.9 \pm 0.5	-0.26 \pm 0.08
JBF036	16 30 54.9845	+00 44 55.134	29.3 \pm 1.5	28.9 \pm 1.5	U	1999 March 05	CnD	44.3 \times 13.5, $-52^{\circ}4$	351	34.7 \pm 1.1	-0.15 \pm 0.05
JBF037	16 30 55.4858	+00 15 38.775	61.9 \pm 3.1	61.4 \pm 3.1	U	1999 March 05	CnD	44.5 \times 13.4, $-52^{\circ}3$	398	20.4 \pm 0.7	+0.89 \pm 0.05
JBF038	16 31 03.5883	+00 21 27.466	9.7 \pm 0.6	9.9 \pm 0.7	U	1999 March 05	CnD	43.0 \times 13.4, $-52^{\circ}1$	425	18.2 \pm 1.0	-0.49 \pm 0.07
JBF039	16 31 15.2280	-00 49 52.827	7.1 \pm 0.6	7.0 \pm 0.6	U	1999 March 05	CnD	41.0 \times 13.4, $-51^{\circ}3$	428	2.9 \pm 0.4	+0.71 \pm 0.13
JBF040	16 31 39.5999	+00 30 41.311	5.2 \pm 0.4	8.4 \pm 0.6	E	1999 March 05	CnD	34.9 \times 13.4, $-50^{\circ}4$	360	4.5 \pm 0.5	+0.50 \pm 0.11
JBF041	16 32 57.7108	-00 33 21.401	199.0 \pm 10.0	203.7 \pm 10.2	E	1999 March 05	CnD	30.9 \times 13.6, $-48^{\circ}1$	715	218.8 \pm 6.6	-0.06 \pm 0.05
JBF042	16 33 07.1572	+00 38 50.622	6.7 \pm 0.5	7.3 \pm 0.5	U	1999 March 05	CnD	29.8 \times 13.6, $-48^{\circ}2$	346	12.9 \pm 0.6	-0.46 \pm 0.07

Table 2 – continued

JBF name	RA (^h ^m ^s)	Dec. ([°] ['] ^{''})	f_{peak} (mJy beam ⁻¹)	S_{int} (mJy)	Morph.	Observation date	Array	Beam size (arcsec ² , PA)	σ_{map} ($\mu\text{Jy beam}^{-1}$)	$S_{1.4}$ (mJy)	$\alpha_{1.4}^{4.86}$
JBF:043	16 55 11.6684	-00 20 19.022	11.3±0.8	11.4±0.8	U	1999 March 02	CnD	58.6 × 4.9, -57:4	496	9.9±0.5	+0.11±0.07
JBF:044	16 56 03.5071	+00 06 10.295	20.9±1.1	21.4±1.2	U	1999 March 02	CnD	48.3 × 4.9, -57:4	452	20.1±1.1	+0.05±0.06
JBF:045	16 56 14.1185	+00 07 36.650	7.2±0.6	7.1±0.6	U	1999 March 02	CnD	47.6 × 4.9, -57:9	439	11.7±0.5	-0.40±0.08
JBF:046	16 56 23.7373	+00 08 29.119	16.4±0.9	17.1±0.9	U	1999 March 02	CnD	44.9 × 4.9, -58:1	462	27.2±1.3	-0.37±0.06
JBF:047	16 56 41.5649	-00 36 04.325	6.3±0.5	6.8±0.5	U	1999 March 02	CnD	42.8 × 5.1, -58:3	385	7.2±0.5	-0.05±0.08
JBF:048	16 57 20.2115	+00 51 29.116	10.0±0.6	10.4±0.7	U	1999 March 02	CnD	35.8 × 5.2, -59:3	390	12.7±0.6	-0.16±0.07
JBF:049	16 57 54.3563	-00 51 37.042	2.6±0.4	6.1±0.4	E	1999 March 02	CnD	32.8 × 5.3, -59:7	393	8.0±0.5	-0.22±0.07
JBF:050	16 58 46.5588	+00 16 17.117	8.5±0.5	8.5±0.5	U	1999 March 02	CnD	26.3 × 5.5, -61:7	324	10.3±0.5	-0.15±0.06
JBF:051	16 58 57.9693	-00 08 51.747	5.5±0.4	8.2±0.5	E	1999 March 02	CnD	25.5 × 5.7, -62:0	325	9.9±0.5	-0.15±0.06
JBF:052	16 59 04.8739	+00 41 01.478	4.5±0.4	5.3±0.4	U	1999 March 02	CnD	24.9 × 5.6, -64:4	334	9.6±0.5	-0.48±0.07
JBF:053	16 59 12.6554	-00 51 29.330	2.3±0.3	7.6±0.5	E	1999 March 02	CnD	24.0 × 5.6, -62:6	272	13.3±1.0	-0.45±0.08
JBF:054	16 59 19.2664	+01 14 19.242	7.1±0.5	7.4±0.5	U	1999 March 02	CnD	22.6 × 5.6, -63:6	298	11.1±0.9	-0.33±0.08
JBF:055	16 59 38.0741	-00 01 03.137	8.2±0.5	8.3±0.5	U	1999 March 02	CnD	21.5 × 5.7, -64:2	299	8.6±0.5	-0.03±0.07
JBF:056	17 00 28.7293	+00 57 44.261	8.2±0.5	8.5±0.5	U	1999 March 02	CnD	18.6 × 6.1, -68:1	310	14.6±0.9	-0.43±0.07
JBF:057	17 01 06.7583	+00 18 49.386	5.6±0.4	6.2±0.4	U	1999 March 02	CnD	17.3 × 6.2, -70:0	310	5.2±0.4	+0.14±0.08
JBF:058	17 01 34.8278	+00 52 22.217	19.0±1.0	19.3±1.0	U	1999 March 02	CnD	16.2 × 6.3, -72:6	281	12.5±0.6	+0.35±0.06
JBF:059	17 02 14.5826	-00 44 40.200	11.9±0.7	17.5±0.9	E	1999 March 02	CnD	15.4 × 6.4, -74:9	280	20.6±0.7	-0.13±0.05
JBF:060	17 02 23.7721	-00 14 03.685	28.2±1.4	28.4±1.4	U	1999 March 02	CnD	15.0 × 6.4, -76:3	291	26.7±0.9	+0.05±0.05
JBF:061	17 02 27.8505	-00 34 39.660	7.7±0.5	8.0±0.5	U	1999 March 02	CnD	14.9 × 6.5, -76:7	274	6.1±0.4	+0.22±0.07
JBF:062	17 03 37.7290	-00 26 43.355	4.8±0.4	5.0±0.4	U	1999 March 02	CnD	14.2 × 6.5, -80:3	265	3.3±0.5	+0.33±0.14
JBF:063	17 03 56.3164	-00 03 14.189	9.3±0.5	7.4±0.5	U	1999 March 02	CnD	14.1 × 6.6, -81:4	268	8.4±0.5	-0.10±0.07
JBF:064	17 04 01.4886	-00 23 17.579	7.7±0.5	13.0±0.7	E	1999 March 02	CnD	14.0 × 6.5, -81:9	277	22.8±0.8	-0.45±0.05
JBF:065	17 04 29.1089	-00 23 14.189	7.7±0.5	8.2±0.5	U	1999 March 02	CnD	13.9 × 6.6, -82:8	297	10.8±0.5	-0.22±0.06
JBF:066	17 04 52.9044	+00 29 49.070	18.0±0.9	17.9±0.9	U	1999 March 02	CnD	13.8 × 6.5, -83:5	288	14.7±1.0	+0.16±0.07
JBF:067	17 30 35.0130	+00 24 38.632	152.6±7.6	160.4±8.0	E	1999 May 21	D	19.8 × 13.0, +5:0	388	213.5±6.4	-0.23±0.05
JBF:068	17 30 50.4616	-00 12 30.348	9.7±0.6	9.8±0.6	U	1999 May 21	D	20.6 × 13.3, +5:1	282	13.6±0.6	-0.26±0.06
JBF:069	19 42 29.4877	+00 39 25.958	4.4±0.4	7.4±0.5	E	1999 May 21	D	20.5 × 13.2, -6:8	291	12.7±0.6	-0.43±0.06
JBF:070	19 42 43.8236	-00 38 16.078	5.9±0.4	5.9±0.4	U	1999 May 21	D	20.7 × 13.2, -6:2	269	5.9±0.5	-0.00±0.09
JBF:071	19 43 00.4706	-00 25 45.225	16.8±0.9	16.9±0.9	U	1999 May 21	D	20.7 × 13.1, -6:1	281	25.0±1.2	-0.31±0.06
JBF:072	19 43 20.0035	-00 44 46.196	13.9±0.8	14.2±0.8	U	1999 May 21	D	20.7 × 13.1, -5:3	299	22.7±0.8	-0.38±0.05
JBF:073	19 43 36.7958	-00 37 41.244	6.3±0.4	6.7±0.4	U	1999 May 21	D	20.7 × 13.1, -3:7	295	11.9±0.5	-0.46±0.06
JBF:074	19 43 48.1149	+00 06 01.453	8.7±0.5	8.9±0.5	U	1999 May 21	D	20.4 × 13.1, -3:7	290	9.8±0.5	-0.08±0.06
JBF:075	19 43 49.4257	+00 22 16.241	6.7±0.4	6.7±0.4	U	1999 May 21	D	20.4 × 13.1, -3:4	269	10.5±0.5	-0.36±0.06
JBF:076	19 44 57.2964	+00 45 47.360	5.0±0.4	5.6±0.4	U	1999 May 21	D	20.2 × 13.1, -0:3	289	3.8±0.4	+0.31±0.10
JBF:077	19 45 22.2084	+00 54 05.237	17.8±0.9	17.9±0.9	U	1999 May 21	D	20.2 × 13.1, -0:4	295	18.2±0.7	-0.01±0.05
JBF:078	19 45 42.1463	+00 17 30.811	22.5±1.2	24.1±1.2	U	1999 May 21	D	20.3 × 13.2, +1:8	279	26.1±0.9	-0.06±0.05
JBF:079	19 46 22.6273	-00 09 08.602	5.9±0.4	6.3±0.4	U	1999 May 21	D	20.5 × 13.3, +4:4	246	8.9±0.5	-0.28±0.07
JBF:080	19 46 45.1455	-00 15 22.942	6.7±0.4	7.3±0.5	U	1999 May 21	D	20.5 × 13.3, +5:5	269	4.8±0.5	+0.34±0.10
JBF:081	19 47 14.2875	+00 27 57.652	5.8±0.4	5.9±0.4	U	1999 May 21	D	20.5 × 13.5, +6:7	273	2.9±0.5	+0.57±0.15
JBF:082	19 47 19.4869	+00 40 09.435	5.0±0.4	7.4±0.5	U	1999 May 21	D	20.4 × 13.5, +7:1	288	11.9±0.6	-0.38±0.07
JBF:083	20 42 46.9679	+00 13 41.761	10.7±0.6	11.5±0.6	U	1999 May 21	D	20.1 × 13.1, +4:0	277	17.3±0.7	-0.33±0.05
JBF:084	20 43 42.1575	+00 01 18.984	44.9±2.3	45.1±2.3	U	1999 May 21	D	20.5 × 13.5, +7:5	314	49.1±1.5	-0.07±0.05
JBF:085	20 44 23.0716	+00 39 12.379	12.7±0.7	12.7±0.7	U	1999 May 21	D	20.9 × 13.5, +8:5	294	10.1±0.5	+0.18±0.06
JBF:086	20 45 16.8655	+00 07 49.853	5.3±0.4	5.8±0.4	U	1999 May 21	D	21.3 × 13.9, +12:0	307	7.7±0.5	-0.23±0.08

Table 2. – continued

JBF name	RA (h m s)	Dec. ($^{\circ}$ ' ")	I_{peak} (mJy beam $^{-1}$)	S_{int} (mJy)	Morph.	Observation date	Array	Beam size (arcsec 2 , PA)	σ_{map} (μ Jy beam $^{-1}$)	$S_{1.4}$ (mJy)	$\alpha_{1.4}^{4.86}$
JBF:087	20 46 51.4309	+00 01 46.361	5.4 \pm 0.4	5.8 \pm 0.4	U	1999 May 21	D	22.0 \times 14.5, +18 $^{\circ}$:1	264	7.9 \pm 0.5	-0.25 \pm 0.08
JBF:088	20 48 18.4294	+00 16 58.972	7.4 \pm 0.5	7.5 \pm 0.5	U	1999 May 21	D	22.6 \times 15.0, +23 $^{\circ}$:3	297	7.0 \pm 0.5	+0.06 \pm 0.08
JBF:089	20 48 42.9746	-00 15 41.124	12.3 \pm 0.7	12.9 \pm 0.7	U	1999 May 21	D	23.0 \times 15.1, +24 $^{\circ}$:4	257	8.6 \pm 0.5	+0.33 \pm 0.06
JBF:090	20 56 23.2526	-00 25 43.302	6.6 \pm 0.4	6.9 \pm 0.4	U	1999 March 05	CnD	19.1 \times 14.4, -20 $^{\circ}$:5	236	12.4 \pm 0.6	-0.47 \pm 0.06
JBF:091	20 56 50.1800	+00 00 24.962	11.8 \pm 0.6	11.9 \pm 0.7	U	1999 March 05	CnD	18.9 \times 14.5, -23 $^{\circ}$:0	268	6.1 \pm 0.5	+0.54 \pm 0.08
JBF:092	20 56 53.3678	+00 10 44.611	10.1 \pm 0.6	10.5 \pm 0.6	U	1999 March 05	CnD	18.8 \times 14.5, -22 $^{\circ}$:8	310	5.7 \pm 0.4	+0.49 \pm 0.07
JBF:093	20 56 53.9599	+00 25 52.058	10.9 \pm 0.6	11.6 \pm 0.6	U	1999 March 05	CnD	18.7 \times 14.5, -22 $^{\circ}$:7	279	14.4 \pm 0.6	-0.17 \pm 0.05
JBF:094	20 56 55.7810	+00 20 32.670	6.5 \pm 0.4	6.7 \pm 0.4	U	1999 March 05	CnD	18.7 \times 14.5, -22 $^{\circ}$:5	254	8.6 \pm 0.5	-0.20 \pm 0.07
JBF:095	20 57 20.3989	+00 12 06.488	68.0 \pm 3.4	69.1 \pm 3.5	U	1999 March 05	CnD	18.6 \times 14.6, -20 $^{\circ}$:6	253	74.9 \pm 2.7	-0.06 \pm 0.05
JBF:096	20 58 09.5097	-00 38 31.454	13.1 \pm 0.7	13.4 \pm 0.7	U	1999 March 05	CnD	18.4 \times 14.7, -9 $^{\circ}$:6	250	19.3 \pm 0.7	-0.29 \pm 0.05
JBF:097	20 59 24.4728	-00 39 28.577	5.8 \pm 0.4	5.9 \pm 0.4	U	1999 March 05	CnD	18.3 \times 14.8, -4 $^{\circ}$:0	234	4.6 \pm 0.6	+0.20 \pm 0.12
JBF:098	20 59 26.5120	+00 26 50.691	30.1 \pm 1.5	30.2 \pm 1.5	U	1999 March 05	CnD	18.1 \times 14.8, -4 $^{\circ}$:0	251	47.2 \pm 1.5	-0.36 \pm 0.05
JBF:099	20 59 38.5741	-00 37 56.224	8.8 \pm 0.5	9.2 \pm 0.5	U	1999 March 05	CnD	18.3 \times 14.9, -2 $^{\circ}$:1	257	14.7 \pm 0.6	-0.38 \pm 0.05
JBF:100	21 00 45.7773	-00 12 26.583	10.3 \pm 0.6	10.3 \pm 0.6	U	1999 March 05	CnD	20.6 \times 14.9, +0 $^{\circ}$:2	255	10.4 \pm 0.5	-0.01 \pm 0.06
JBF:101	21 01 28.2119	+00 19 50.116	15.3 \pm 0.8	18.5 \pm 1.0	E	1999 March 05	CnD	20.1 \times 14.9, +4 $^{\circ}$:2	279	31.4 \pm 1.3	-0.43 \pm 0.07
JBF:102	21 02 19.8957	+00 40 27.380	5.9 \pm 0.4	6.2 \pm 0.4	U	1999 March 05	CnD	20.1 \times 15.2, +7 $^{\circ}$:9	284	6.6 \pm 0.4	-0.05 \pm 0.05
JBF:103	21 02 20.0844	+00 29 52.728	13.4 \pm 0.7	13.6 \pm 0.7	U	1999 March 05	CnD	20.2 \times 15.2, +8 $^{\circ}$:2	283	20.2 \pm 0.7	-0.32 \pm 0.05
JBF:104	21 41 38.5314	+00 03 19.690	5.7 \pm 0.3	6.3 \pm 0.4	U	1999 May 21	D	21.4 \times 13.9, +12 $^{\circ}$:4	188	10.5 \pm 0.5	-0.41 \pm 0.06
JBF:105	21 43 24.3029	+00 35 02.093	61.8 \pm 3.1	61.6 \pm 3.1	U	1999 May 21	D	21.9 \times 14.6, +18 $^{\circ}$:5	361	45.1 \pm 1.4	+0.25 \pm 0.05
JBF:106	21 44 06.2709	-00 28 57.672	23.2 \pm 1.2	23.6 \pm 1.2	U	1999 May 21	D	23.0 \times 15.0, +24 $^{\circ}$:1	367	38.2 \pm 1.5	-0.39 \pm 0.05
JBF:107	21 44 11.6640	+00 03 21.485	7.4 \pm 0.5	7.5 \pm 0.5	U	1999 May 21	D	23.2 \times 15.3, +26 $^{\circ}$:6	284	12.1 \pm 0.6	-0.38 \pm 0.07
JBF:108	21 44 19.9069	+00 20 54.962	11.3 \pm 0.9	16.5 \pm 1.1	U	1999 May 21	D	23.2 \times 15.4, +27 $^{\circ}$:4	437	24.0 \pm 1.3	-0.30 \pm 0.07
JBF:109	21 44 29.4053	+00 37 23.859	7.5 \pm 0.5	7.5 \pm 0.5	E	1999 May 21	D	24.2 \times 15.3, +33 $^{\circ}$:1	313	6.8 \pm 0.5	+0.08 \pm 0.09
JBF:110	21 45 35.8249	-00 03 35.724	5.1 \pm 0.4	5.8 \pm 0.4	U	1999 May 21	D	23.7 \times 14.7, +36 $^{\circ}$:1	324	6.0 \pm 0.5	-0.03 \pm 0.09
JBF:111	21 46 13.3517	+00 09 31.065	9.8 \pm 0.6	10.2 \pm 0.6	U	1999 May 21	D	23.9 \times 14.8, +37 $^{\circ}$:4	374	7.3 \pm 0.5	+0.27 \pm 0.07
JBF:112	21 46 43.0390	+00 31 53.336	12.6 \pm 0.7	12.6 \pm 0.7	U	1999 May 21	D	25.4 \times 14.9, +39 $^{\circ}$:2	370	9.2 \pm 0.5	+0.25 \pm 0.06
JBF:113	22 13 16.1962	-00 34 10.244	15.3 \pm 0.9	15.5 \pm 0.9	U	1999 May 21	D	27.7 \times 15.1, +43 $^{\circ}$:5	394	23.8 \pm 0.8	-0.34 \pm 0.05
JBF:114	22 14 33.1704	-00 16 22.141	6.6 \pm 0.5	7.3 \pm 0.5	U	1999 May 21	D	24.7 \times 14.9, +38 $^{\circ}$:8	327	13.3 \pm 0.6	-0.48 \pm 0.06
JBF:115	23 13 09.7213	+00 08 03.176	6.6 \pm 0.5	7.0 \pm 0.5	U	1999 May 21	D	24.7 \times 14.9, +38 $^{\circ}$:8	407	8.7 \pm 0.6	-0.17 \pm 0.08
JBF:116	23 13 22.6162	-00 02 16.154	10.0 \pm 0.6	10.4 \pm 0.6	U	1999 May 21	D	24.7 \times 14.9, +38 $^{\circ}$:8	389	7.1 \pm 0.5	+0.31 \pm 0.07
JBF:117	23 15 48.1771	+00 07 22.011	7.0 \pm 0.5	10.3 \pm 0.6	E	1999 May 21	D	28.5 \times 15.2, +44 $^{\circ}$:9	358	11.1 \pm 1.2	-0.06 \pm 0.10

Table 3. The JBF 8.46 GHz VLA data. Columns 1–5 are the same as in Table 2. The rms noise in of each map is given in column 6. The 1.4–8.46 GHz and 4.86–8.46 GHz spectral indices of each source are given in Columns 7 and 8, respectively.

JBF Name	RA (^h ^m ^s)	Dec. ([°] ['] ^{''})	I_{peak} (mJy beam ⁻¹)	S_{int} (mJy)	σ_{map} (μ Jy beam ⁻¹)	$\alpha_{1.4}^{8.46}$	$\alpha_{4.86}^{8.46}$
JBF.001	03 57 51.5304	+00 30 48.065	5.4±0.3	5.3±0.3	188	-0.14±0.05	-0.72±0.15
JBF.009	07 59 20.6127	-00 14 02.567	36.6±1.8	36.7±1.9	227	-0.27±0.03	-0.20±0.13
JBF.011	07 59 54.0308	+00 05 09.092	14.5±0.7	14.4±0.7	167	-0.16±0.03	-0.06±0.13
JBF.012	08 01 18.8764	+00 20 12.921	4.7±0.3	4.6±0.3	174	-0.09±0.06	-0.60±0.16
JBF.014	13 14 33.0024	+00 06 37.250	7.2±0.4	7.4±0.4	209	+0.09±0.05	-0.27±0.14
JBF.016	13 30 20.1733	+00 17 18.880	9.5±0.5	9.4±0.5	210	-0.24±0.04	-0.57±0.14
JBF.019	14 14 16.5521	+00 24 06.398	13.4±0.7	13.8±0.7	236	-0.19±0.04	-0.42±0.13
JBF.023	14 27 53.7814	+00 00 38.970	15.3±0.8	15.8±0.8	231	-0.21±0.03	-0.10±0.13
JBF.028	14 29 56.6346	-00 17 56.028	13.7±0.8	13.6±0.7	235	-0.04±0.04	-0.82±0.13
JBF.029	14 30 31.3066	-00 09 07.630	61.5±3.1	61.5±3.1	247	+0.03±0.03	+0.21±0.13
JBF.031	15 00 27.8334	+00 02 24.648	15.4±0.8	15.7±0.8	241	-0.17±0.04	-0.68±0.13
JBF.033	16 29 56.7259	+01 01 40.862	70.4±3.5	70.1±3.5	305	+0.25±0.03	-0.52±0.13
JBF.036	16 30 54.9439	+00 44 55.626	24.2±1.2	24.3±1.2	223	-0.20±0.03	-0.31±0.13
JBF.037	16 30 55.4448	+00 15 38.486	55.2±2.8	54.9±2.8	278	+0.55±0.03	-0.20±0.13
JBF.039	16 31 15.1854	-00 49 53.058	4.1±0.3	4.2±0.3	224	+0.21±0.09	-0.92±0.20
JBF.041	16 32 57.6775	-00 33 21.107	138.0±6.9	151.8±7.6	619	-0.20±0.03	-0.53±0.13
JBF.043	16 55 11.5294	-00 20 17.940	10.1±0.5	10.1±0.6	217	+0.01±0.04	-0.22±0.17
JBF.048	16 57 20.1660	+00 51 29.720	11.7±0.6	11.7±0.6	200	-0.05±0.04	+0.21±0.15
JBF.050	16 58 46.5365	+00 16 17.281	6.4±0.4	6.7±0.4	175	-0.24±0.04	-0.43±0.15
JBF.055	16 59 38.0688	-00 01 03.064	8.0±0.5	7.8±0.4	211	-0.05±0.04	-0.11±0.14
JBF.058	17 01 34.8165	+00 52 22.300	14.2±0.7	14.4±0.8	213	+0.08±0.04	-0.53±0.14
JBF.059	17 02 14.5941	-00 44 40.820	6.7±0.4	6.9±0.4	261	-0.61±0.04	-1.68±0.14
JBF.060	17 02 23.7645	-00 14 03.749	21.8±1.1	21.7±1.1	186	-0.12±0.03	-0.49±0.13
JBF.061	17 02 27.8333	-00 34 39.740	5.9±0.4	5.8±0.4	236	-0.03±0.05	-0.58±0.17
JBF.063	17 03 56.3039	-00 36 08.246	5.2±0.3	5.4±0.3	193	-0.25±0.05	-0.57±0.16
JBF.065	17 04 29.0713	-00 23 17.480	5.1±0.3	5.1±0.3	206	-0.42±0.04	-0.86±0.15
JBF.066	17 04 52.9034	+00 29 48.961	14.8±0.8	14.8±0.8	228	+0.00±0.05	-0.34±0.13
JBF.067	17 30 35.0102	+00 24 38.680	148.2±7.4	149.5±7.5	453	-0.20±0.03	-0.13±0.13
JBF.068	17 30 50.4911	-00 12 30.220	7.3±0.4	7.2±0.4	204	-0.35±0.04	-0.56±0.15
JBF.071	19 43 00.8781	-00 25 42.166	17.2±0.9	18.2±0.9	163	-0.18±0.04	+0.13±0.13
JBF.072	19 43 20.0212	-00 44 46.298	10.9±0.6	10.9±0.6	151	-0.41±0.04	-0.48±0.14
JBF.074	19 43 48.5240	+00 06 03.718	7.3±0.5	7.6±0.5	399	-0.14±0.05	-0.28±0.16
JBF.075	19 43 49.8227	+00 22 19.418	5.5±0.5	5.5±0.5	432	-0.36±0.06	-0.36±0.20
JBF.077	19 45 22.1944	+00 54 05.263	13.2±0.8	13.6±0.8	392	-0.16±0.04	-0.50±0.14
JBF.078	19 45 42.1442	+00 17 31.461	18.4±1.0	18.4±1.0	473	-0.19±0.04	-0.49±0.13
JBF.079	19 46 22.6260	-00 09 07.860	7.1±0.5	7.4±0.5	404	-0.10±0.05	+0.29±0.17
JBF.080	19 46 45.1625	-00 15 23.079	9.8±0.7	10.3±0.7	489	+0.42±0.07	+0.62±0.17
JBF.081	19 47 14.3040	+00 27 57.608	4.6±0.3	4.8±0.3	174	+0.28±0.10	-0.37±0.17
JBF.083	20 42 46.9871	+00 13 41.642	6.5±0.4	7.0±0.4	146	-0.50±0.04	-0.90±0.14
JBF.084	20 43 42.1651	+00 01 18.864	45.7±2.3	45.7±2.3	172	-0.04±0.03	+0.02±0.13
JBF.085	20 44 23.0617	+00 39 12.325	14.0±0.7	13.8±0.7	147	+0.17±0.04	+0.15±0.14
JBF.088	20 48 18.4187	+00 16 59.043	9.1±0.5	9.6±0.5	173	+0.18±0.05	+0.45±0.15
JBF.089	20 48 42.9790	-00 15 41.360	11.0±0.6	11.1±0.6	141	+0.14±0.04	-0.27±0.14
JBF.091	20 56 50.1720	+00 00 25.320	10.3±0.5	10.4±0.6	173	+0.30±0.06	-0.24±0.15
JBF.092	20 56 53.3042	+00 10 44.908	13.1±0.7	13.2±0.7	161	+0.47±0.05	+0.41±0.14
JBF.093	20 56 53.9115	+00 25 51.480	9.1±0.5	9.0±0.5	162	-0.26±0.04	-0.46±0.14
JBF.094	20 56 55.7629	+00 20 32.661	5.0±0.3	5.1±0.3	143	-0.29±0.05	-0.49±0.15
JBF.095	20 57 20.3820	+00 12 07.314	52.1±2.6	52.1±2.6	197	-0.20±0.03	-0.51±0.13
JBF.098	20 59 26.5265	+00 26 51.400	22.0±1.1	22.2±1.1	143	-0.42±0.03	-0.56±0.13
JBF.100	21 00 45.7638	-00 12 27.916	6.9±0.4	6.9±0.4	182	-0.23±0.04	-0.72±0.15
JBF.103	21 02 20.1050	+00 29 52.447	9.1±0.5	9.3±0.5	150	-0.43±0.04	-0.69±0.13
JBF.105	21 43 24.3565	+00 35 02.778	71.4±3.6	71.6±3.6	207	+0.26±0.03	+0.27±0.13
JBF.108	21 44 19.8702	+00 20 55.712	8.5±0.5	8.5±0.5	221	-0.58±0.04	-1.20±0.16
JBF.109	21 44 29.2193	+00 37 22.842	6.1±0.3	6.1±0.4	164	-0.06±0.05	-0.37±0.17
JBF.111	21 46 13.3127	+00 09 30.800	10.0±0.5	10.0±0.5	147	+0.17±0.05	-0.04±0.14
JBF.112	21 46 42.9687	+00 31 53.780	13.7±0.7	13.6±0.7	142	+0.22±0.04	+0.14±0.14
JBF.113	22 13 16.2011	-00 34 10.979	12.1±0.6	12.5±0.6	155	-0.36±0.03	-0.39±0.14
JBF.115	23 13 09.8539	+00 08 05.535	6.4±0.4	6.5±0.4	161	-0.16±0.05	-0.13±0.17
JBF.116	23 13 22.7329	-00 02 12.944	8.2±0.4	7.9±0.4	184	+0.06±0.05	-0.50±0.14

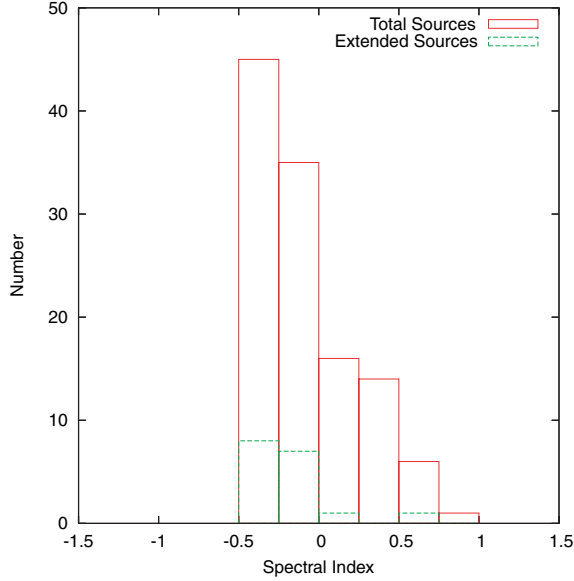


Figure 1. The spectral index distribution of the JBF sample measured between 1.4 and 4.86 GHz. The solid (red) line represents the distribution for the total sample, whereas the broken (green) line is the distribution for the extended sources in JBF.

extended source fraction of the full JBF sample given above). The 33 sources which make up the FIRST subsample are JBF.013–JBF.031 and JBF.104–JBF.117. We define sources as unresolved at 1.4 GHz if they consist of a single radio component with a deconvolved FWHM of less than 4 arcsec within 30 arcsec of the JBF position in the FIRST radio maps. Note that during the selection process of the JBF sample we used a search radius of 70 arcsec in order to remain consistent with the selection process used by CLASS. Here, we only consider radio emission within 30 arcsec of the JBF position because we are now looking for evidence for jet emission associated with each JBF source. Using the above criteria, we find that only three JBF sources (JBF.025, JBF.026 and JBF.031) show signs of extension in the FIRST radio maps. These three sources were also identified as extended by the 4.86-GHz pseudo-survey observations. The three other extended sources from the 4.86-GHz pseudo-survey

imaging (JBF.020, JBF.108 and JBF.117) had compact structures in the FIRST maps, but were found to have another independent radio source between 30 and 70 arcsec from the JBF position. The FIRST images of JBF.025, JBF.026 and JBF.031 are given in Fig. 2 and a brief description of each source is given below.

JBF.025 appears as a single extended radio source with a FIRST 1.4-GHz flux density of 7.6 mJy and a deconvolved FWHM of 4.67 arcsec. The radio structure appears unremarkable with a slight extension to the north. There is another FIRST radio source ~ 45 arcsec towards the east.

JBF.026 shows clear extended structure elongated towards the south-west. The 1.4-GHz flux density measured by the FIRST survey is 12.1 mJy and the deconvolved FWHM is 7.97 arcsec

JBF.031 has the most interesting radio structure of the three extended JBF sources. JBF.031 consists of three radio components extending in a north–south direction separated by 27.5 arcsec. The most southern component, JBF.031a, has the largest 1.4-GHz flux density of the three radio components (12.3 mJy) and is the most compact (deconvolved FWHM is 1.28 arcsec). Also, JBF.031a is the only radio component to be detected at 8.46 GHz during the pseudo-survey observations (see Table 3). The spectral index of JBF.031a between 1.4 (FIRST) and 8.46 GHz (JBF) is flat/rising ($\alpha_{1.4}^{8.46} = +0.13 \pm 0.06$). Therefore, we associate JBF.031a as the radio core of JBF.031. The remaining two components to the north, JBF.031b and JBF.031c, have 1.4-GHz flux densities of 3.5 and 5.1 mJy and deconvolved sizes of 4.57 and 4.96 arcsec, respectively. Both JBF.031b and JBF.031c have structures consistent with a radio jet.

Assuming that the FIRST subsample is representative of the whole JBF sample, we find that only 9 per cent of the JBF sample show evidence for extended jet emission, with the vast majority being unresolved and compact. Of course, further 1.4-GHz imaging of the remaining 84 JBF sources not observed by FIRST could confirm this result. In general, we find that the JBF sample is composed of compact radio sources with little or no evidence of extended jet emission on the arcsec scales probed here.

4.2 Radio source number counts

The differential number counts of the CLASS parent population have been determined by combining the JBF and CLASS complete

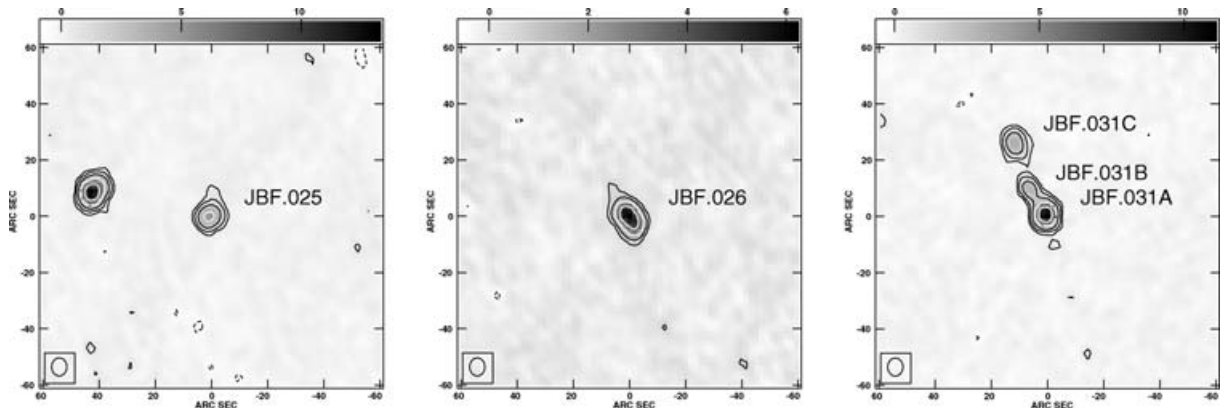


Figure 2. The FIRST 1.4-GHz radio maps (Becker et al. 1995) of the extended radio sources from the JBF sample. Left: JBF.025 shows a slight extension to the north and another (possibly independent) radio source 45 arcsec to the east. The contour levels are $(-3, 3, 6, 12, 24, 48) \times 170 \mu\text{Jy beam}^{-1}$. Centre: JBF.026 shows extension towards the south-west. The contour levels are $(-3, 3, 6, 12, 24) \times 157 \mu\text{Jy beam}^{-1}$. Right: JBF.031 consists of three radio components extending to the north; a core (A) and two jet components (B and C). The contour levels are $(-3, 3, 6, 12, 24, 48) \times 142 \mu\text{Jy beam}^{-1}$. The grey-scales in each map are in units of mJy beam^{-1} .

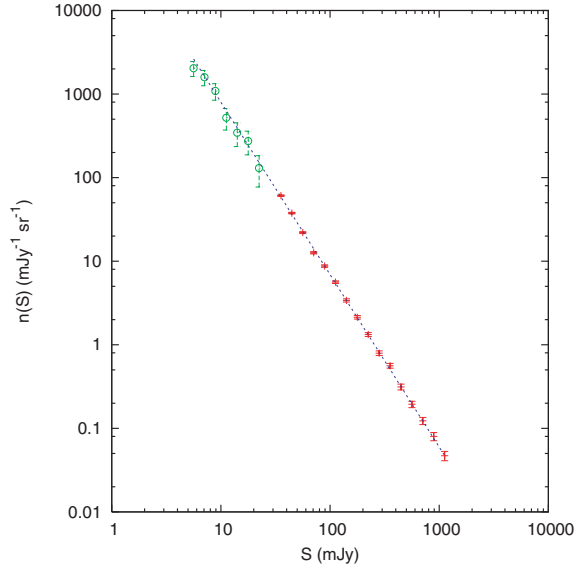


Figure 3. The CLASS parent population differential number counts. The data above 30 mJy (red crosses) are taken from CLASS and the data below 30 mJy (green circles) are taken from JBF. The dashed line is the best-fitting power law, which has an index of $\eta = -2.06 \pm 0.01$.

samples. We excluded from our analysis the number counts data from the JBF sample at $S \gtrsim 25$ mJy because (i) the small number of JBF sources with flux densities above 25 mJy led to large uncertainties in the number counts per flux density bin (60–100 per cent), and (ii) the CLASS complete sample provides excellent number counts information over the 30 mJy to ~ 1 Jy flux density range. Fig. 3 shows the differential number counts of flat-spectrum radio sources as a function of flux density. The JBF number counts follow on smoothly from those obtained with the CLASS complete sample. Using a least-squares fitting technique, we find the differential number counts of flat-spectrum radio sources with $S_{4.85} \geq 5$ mJy are described by the power law

$$n(S) = (6.91 \pm 0.42) \left(\frac{S_{4.85}}{100 \text{ mJy}} \right)^{-2.06 \pm 0.01} \text{ mJy}^{-1} \text{ sr}^{-1}. \quad (1)$$

The reduced χ^2 of the fit is 1.31 and the number of degrees of freedom (ndf) is 21. Clearly, this power-law fit has been heavily weighted by the CLASS complete sample data which have very small uncertainties in the number of sources per flux density bin. As the CLASS gravitational lensing statistics will be particularly sensitive to any change in the differential number counts slope, η , where $n(S) \propto S^\eta$, below 30 mJy, two separate power laws have been fitted to the parent population data above and below the CLASS 30-mJy flux density limit. We find from the resulting least-squares fits,

$$n(S) = (7.97 \pm 2.23) \left(\frac{S_{4.85}}{100 \text{ mJy}} \right)^{-1.96 \pm 0.12} \text{ mJy}^{-1} \text{ sr}^{-1}, \quad (2)$$

for $5 \leq S < 30$ mJy (reduced $\chi^2 = 0.31$; ndf = 5) and

$$n(S) = (6.85 \pm 0.50) \left(\frac{S_{4.85}}{100 \text{ mJy}} \right)^{-2.07 \pm 0.02} \text{ mJy}^{-1} \text{ sr}^{-1}, \quad (3)$$

for $S \geq 30$ mJy (reduced $\chi^2 = 1.73$; ndf = 14). The large uncertainty in the slope below 30 mJy is due to the small number of sources in the JBF sample.

The differential number counts slope below 30 mJy presented here is slightly different to the result reported by Chae et al. (2002) ($\eta = -1.97 \pm 0.14$). The small change in η below 30 mJy is due to a recent update of the NVSS catalogue in 2004 which led to an increase in the number of flat-spectrum radio sources within the JBF sample. This change in η has a negligible effect on the CLASS gravitational lensing statistics, with Ω_Λ unchanged from the result published by Chae et al. (2002).

4.3 Fraction of radio sources with flat radio spectra

In Fig. 4, the percentage of radio sources with flat radio spectra ($\alpha_{1.4}^{4.85} \geq -0.5$) as a function of flux density is presented. Those data above 30 mJy come from the combination of the NVSS and GB6 catalogues, and those data below 30 mJy are taken from the 4.86-GHz pseudo-survey. There is a clear change in the spectral composition of the radio source population with flux density. At high flux densities (> 1 Jy), the radio source population is dominated by the powerful flat-spectrum quasars. As the quasar population declines with flux density (e.g. Falco et al. 1998; Marlow et al. 2000; Muñoz et al. 2003), so does the fraction of sources with flat radio spectra. From ~ 10 to 100 mJy, the fraction remains constant with about one in four radio sources having flat spectra. Also, those data from the pseudo-survey appear to closely match the results from NVSS and GB6 at the transition point around 30 mJy, although the uncertainties in the fraction of sources with flat spectra from the pseudo-survey are quite large. Interestingly, there is a hint of an increase in the fraction of radio sources with flat radio spectra below 10 mJy to about one in three radio sources. A possible explanation for this increase is that the pseudo-survey observations partially or completely resolved out extended steep-spectrum radio sources which would have otherwise been detected by the ~ 3.5 arcmin beam of the GB6 survey. Although this does not affect the number of compact flat-spectrum radio sources found by the VLA pseudo-survey, it could result in an increase in the fraction of radio sources identified with

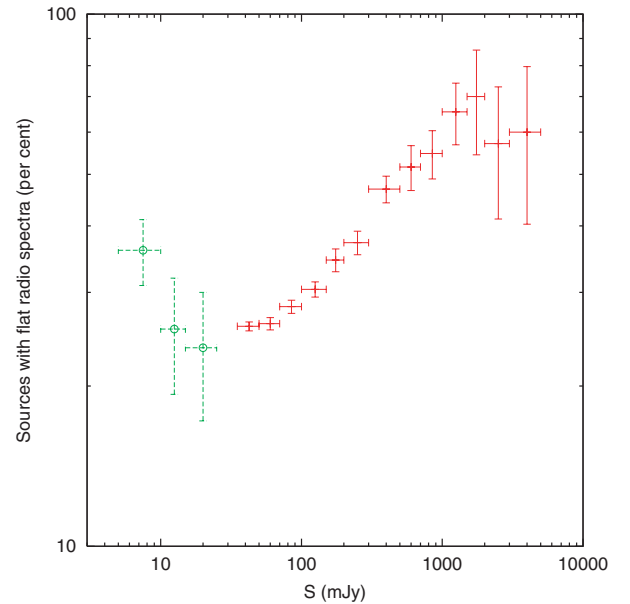


Figure 4. The percentage of radio sources with flat radio spectra at 4.85 GHz as a function of flux density. The data above 30 mJy (red crosses) have been calculated using the NVSS and GB6 catalogues. The data below 30 mJy (green circles) have been taken from the VLA pseudo-survey.

flat spectra at the survey limit (~ 5 mJy). Alternatively, the fraction of radio sources with flat radio spectra may be genuinely increasing. However, a much larger survey of the mJy level radio source population using a radio array/telescope with a greater sensitivity to extended emission will need to be carried out to confirm this intriguing result.

5 CONCLUSIONS

The selection of the JBF sample from a 4.86-GHz VLA pseudo-survey has been presented. We find the vast majority of the 117 flat-spectrum radio sources within JBF to be compact and unresolved over the arcsec scales probed here. Using the JBF and CLASS complete samples, we have determined the differential number counts slope of the CLASS parent population above and below 30 mJy to be -2.07 ± 0.02 and -1.96 ± 0.12 , respectively. The parent population number counts information presented here forms a vital part of the CLASS gravitational lensing statistics.

However, these number counts must be coupled with complete redshift information for the JBF sample because the lensing optical depth is strongly dependent on the redshift of the background source (e.g. Turner et al. 1984). The analysis of the CLASS gravitational lensing statistics performed by Chae et al. (2002) assumed that the mean redshift of the flat-spectrum radio source population below 25 mJy was $\bar{z} = 1.27$, i.e. the same as for brighter samples of flat-spectrum radio sources (e.g. Marlow et al. 2000). If the true mean redshift of the flat-spectrum radio source population below 25 mJy differs from 1.27 by ± 0.1 , this would result in a change of ∓ 0.06 in the value of Ω_{Λ} obtained from the CLASS gravitational lensing statistics (see fig. 10 of Chae 2003). As such, it is crucial we establish the redshift distribution of faint flat-spectrum radio sources below the CLASS flux density limit. In a companion paper to this one (McKean et al., in preparation), we will present the optical/infrared followup of a small subsample of JBF sources with flux densities between 5 and 15 mJy. Our preliminary results, based on a combination of redshifts obtained from spectroscopy and photometry, suggest that the mean redshift of the JBF selected subsample is $\bar{z} \sim 1.2$. Therefore, we expect little change in the value of Ω_{Λ} once the redshift information for the parent population below 25 mJy is incorporated into the CLASS gravitational lensing statistics analysis.

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