



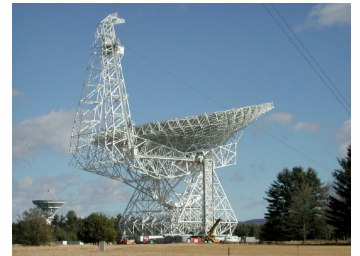
Radio Astronomy

Lecture II

The Future of Radio Astronomy

Lecturer: Michael Wise (wise@astron.nl)

May 29th, 2017



This lecture we will discuss the future of radio astronomy, both opportunities and challenges. We'll talk about a wide range of new instruments that are being constructed but also the changes these new telescopes will require to the way we do radio astronomy.

We'll finish up with a look further ahead to some more speculative directions the field may take.

Outline

- The Square Kilometre Array (SKA)
- SKA Pathfinders
- SKA Computational Challenges
- Data Intensive Astronomy
- Beyond the SKA



The Square Kilometre Array

The Square Kilometre Array



- Next generation global radio telescope
- Frequency range: 0.1–25 GHz, 100x as sensitive as JVLA
- Being constructed in South Africa and Western AU
- Online in 2021 and expected to cost €1.5 Billion

The SKA is project that is already underway. It is a truly global project involving countries from all around the world including Europe. It will be the most powerful radio telescope on the plant in the coming decade.

The Netherlands is a major leader in the international consortium to design and build the SKA.

Current facilities have reached the point at which single nations can still fund them.

To make a big enough step forward in scientific capability, we must cooperate globally.

“The Hydrogen Array”

199 JASPC...19...428W

428

*Radio Interferometry: Theory, Techniques and Applications,
IAU Coll. 131, ASP Conference Series, Vol. 19, 1991,
T.J. Cornwell and R.A. Perley (eds.)*

THE HYDROGEN ARRAY

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ABSTRACT The time is ripe for planning an array with a collecting area of 1 km^2 (14 times larger than Arecibo and 75 times larger than the VLA). In view of its major astronomical target I have dubbed this concept ‘The Hydrogen Array’, although $1 \mu\text{Jy}$ continuum sources will also be reliably detected. I present some initial thoughts about the issues involved.

INTRODUCTION

Since the late 1960s radioastronomers have increased the capability of their instruments many fold. The maximum resolution achieved with interferometry has increased from ~ 50 milliarcsec to ~ 50 microarcsec; the highest frequency in use has gone from $\sim 10 \text{ GHz}$ to $> 350 \text{ GHz}$ and the aperture plane coverage has improved from that of the One-Mile Telescope to that of the multi-configuration VLA. However, in terms of raw sensitivity the improvement has been less dramatic. The Arecibo telescope remains the world’s largest and the improvements to system noise temperatures at decimetric and centimetric wavelengths have been relatively small (≤ 5). Despite its limitations in sky and frequency coverage, the scientific output of the Arecibo telescope amply demonstrates the advantage of a collecting area 5–10 times larger than that of

- Originally motivated by high redshift HI studies
- Detect 21 cm hydrogen emission line (“H I”) from normal galaxies anywhere in Universe ($z \sim 2$)
- Current science case is much broader



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Radio Astronomy - 5214RAAS6Y

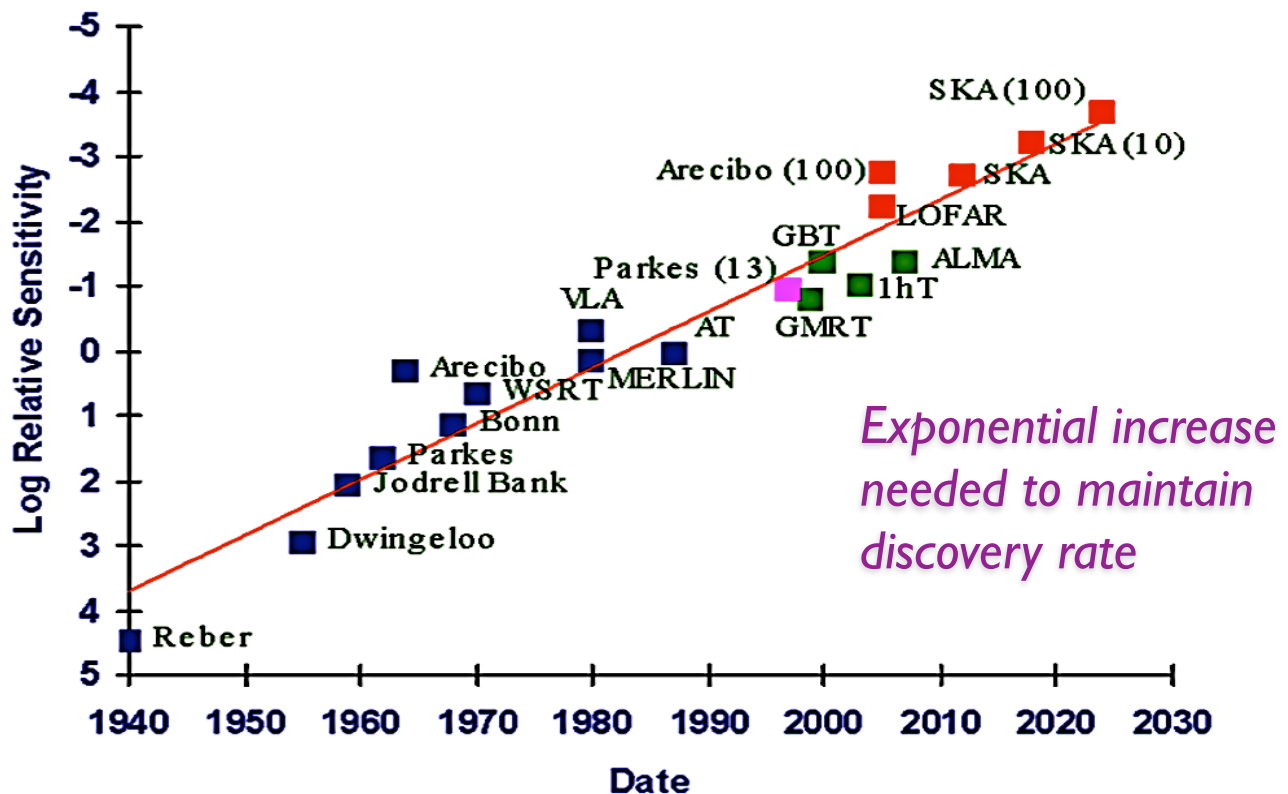
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The SKA has originally conceived as a way to search for highly redshifted HI.

The name comes from the collecting area necessary to detect HI in a normal galaxy at a redshift of 2.

The science case now includes almost every topic in radio astronomy, but there are a set of priorities.

Radio Telescope Sensitivity

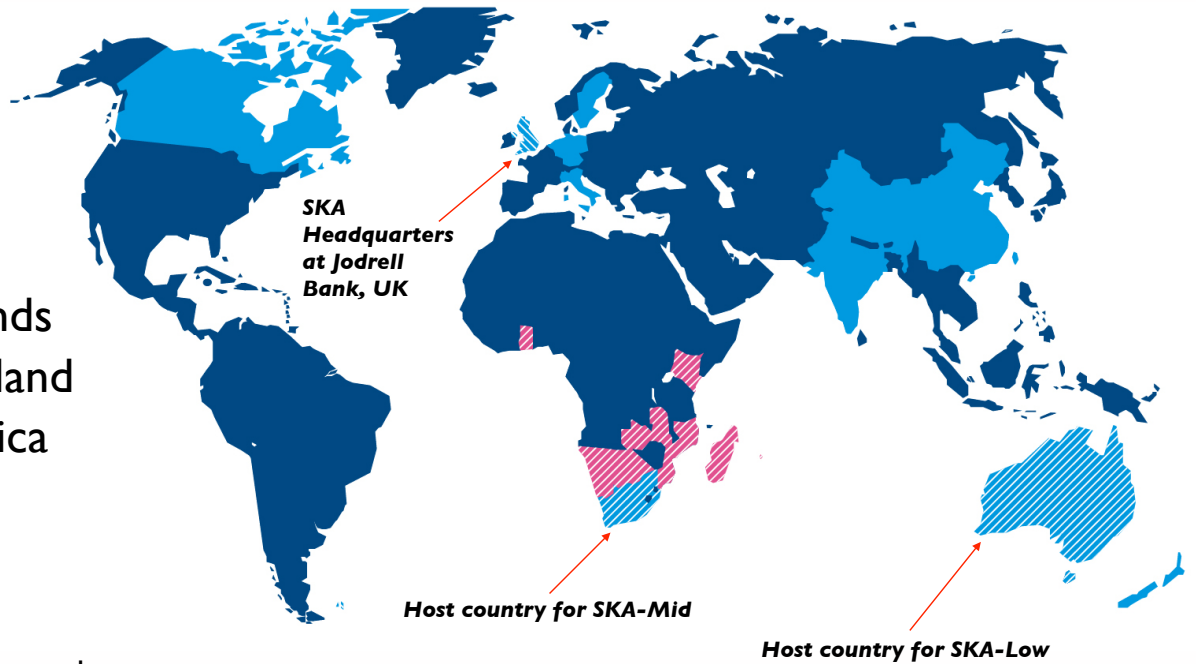


The plot shows the rate at which the sensitivity of radio telescopes has been increasing over time. If we want to keep making new discoveries, we have to keep climbing this curve. Since 2000, the new more sensitive telescopes have been international collaborations. Some of the increase is due to size, but in many cases in new technology and computational techniques. We will need steady improvements in all three aspects if we want to stay on this curve.

The Square Kilometre Array

- Australia
- Canada
- China
- India
- Italy
- Netherlands
- New Zealand
- South Africa
- Sweden
- UK

Potential new members:
Spain, Portugal, Germany,
France, others...



Formation and Evolution of Galaxies • The Dawn of Galaxies: Searching for the Epoch of First Light • 21-cm Emission and Absorption Mechanisms • Preheating the IGM • **SKA Imaging of Cosmological HI** • Large Scale Structure and Galaxy Evolution • A Deep SKA HI Pencil Beam Survey • Large scale structure studies from a shallow, wide area survey • The Ly- α forest seen in the 21-cm HI line • **High Redshift CO** • Deep Continuum Fields • Extragalactic Radio Sources • The SubmicroJansky Sky • Probing Dark Matter with **Gravitational Lensing** • Activity in Galactic Nuclei • The SKA and Active Galactic Nuclei • Sensitivity of the SKA in VLBI Arrays • Circumnuclear MegaMasers • H₂O megamasers • OH Megamasers • Formaldehyde Megamasers • The **Starburst** Phenomenon • **Interstellar Processes** • HII Regions: High Resolution Imaging of Thermal Emission • Centimetre Wavelength Molecular Probes of the ISM • **Supernova Remnants** • The Origin of Cosmic Rays • Interstellar Plasma Turbulence • Recombination Lines • Magnetic Fields • Rotation Measure Synthesis • Polarization Studies of the Interstellar Medium in the Galaxy and Nearby External Galaxies • Formation and Evolution of Stars • Continuum Radio Emission from Stars • Imaging the Surface of Stars • Red Giants and Supergiant Stars • Star Formation • Protostellar Cores • Protostellar Jets • Uncovering the Evolutionary Sequence • Magnetic Fields in Protostellar Objects • Cool Star Astronomy • The Radio Sun • Observing Solar Analogs at Radio Wavelengths • Where are the many other Radio Suns? • Flares and Microflares • X-ray Binaries • Relativistic Electrons from X-ray Transients • The Faint Persistent Population • Imaging of Circumstellar Phenomena • Stellar Astrometry • Supernovae • Radio **Supernovae** • The Radio After-Glows of **Gamma-ray Bursts** • Pulsars • Pulsar Searches • **Pulsar Timing** • Radio Pulsar Timing and General Relativity • Solar System Science • Thermal Emission from Small Solar System Bodies • Asteroids • Planetary Satellites • **Kuiper Belt Objects** • Radar Imaging of Near Earth Asteroids • The Atmosphere and Magnetosphere of Jupiter • Comet Studies • Solar Radar • Coronal Scattering • Formation and Evolution of Life • Detection of **Extrasolar Planets** • Pre-Biotic Interstellar Chemistry • The Search for **Extraterrestrial Intelligence**

SKA Science

The science case for the SKA is very broad.
 Anything you can do with a radio telescope,
 someone will do with the SKA.
 As we have learned, we often must tailor our
 telescope to the science we want to do.
 Put another way, the science we want to do defines
 what our telescope should look like.
 The SKA will be good for many types of science, but it is
 been optimized for a few.

The Science Case

The First Stars

Cosmic Evolution

Cosmic Magnetism

Gravitational Physics

Origins of Life

The science case for the SKA is organized around five central themes.

Cosmology and the epoch of reionization, early in the evolution of the universe when the first stars formed.

The growth of cosmic structure including the formation and evolution of galaxies and black holes. Tests of fundamental physics including how magnetic fields are formed and detection of gravity waves.

The study of exoplanets and planet formation and searches for biomarkers.

SKA Specifications

- Frequency range: **0.1-25 GHz**
- Bandwidth: **0.3 -10 GHz**
- Sensitivity ($A_{\text{eff}} / T_{\text{sys}}$): **$2 \times 10^4 \text{ m}^2 \text{ K}^{-1}$**
- Baseline distribution: 50% within 5 km, some baselines beyond 3000 km
- Imaging Field OfView: **1° @ 1.4 GHz**
- Angular Resolution: **$0.1''$ @ 1.4 GHz**
- Image Dynamic Range: **10^6 @ 1.4 GHz**
- SKA 1: **300 dishes, 50 AA fields**
- SKA 2: **2700 dishes, 250 AA fields**
- Estimated cost: $\sim \text{€}1.5$ billion
- Operations in **2019** (SKA1), **2024** (SKA2)

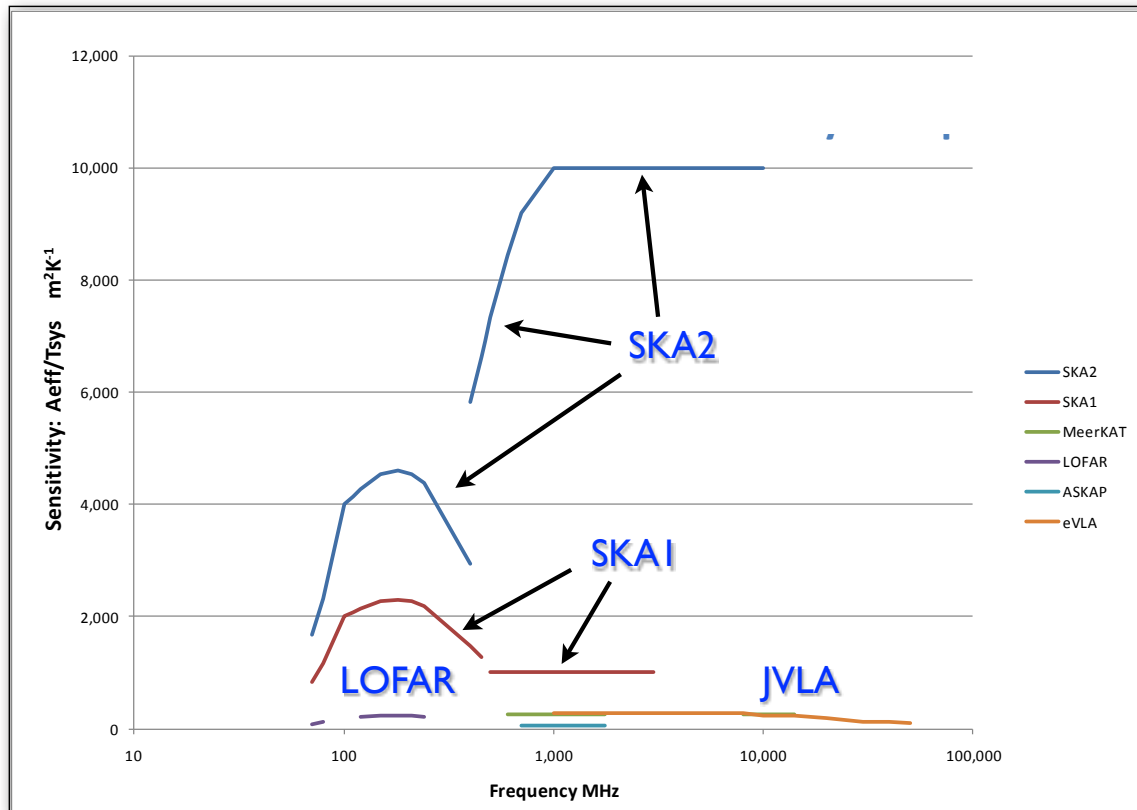


The SKA will cover a wide range in frequencies. To do so, it will actually consist of several distinct telescopes.

It will be built in two stages. SKA1 will consist of roughly 10% of the fully envisioned SKA.

Construction is expected to begin in 2017 with science operations starting before the end of the decade.

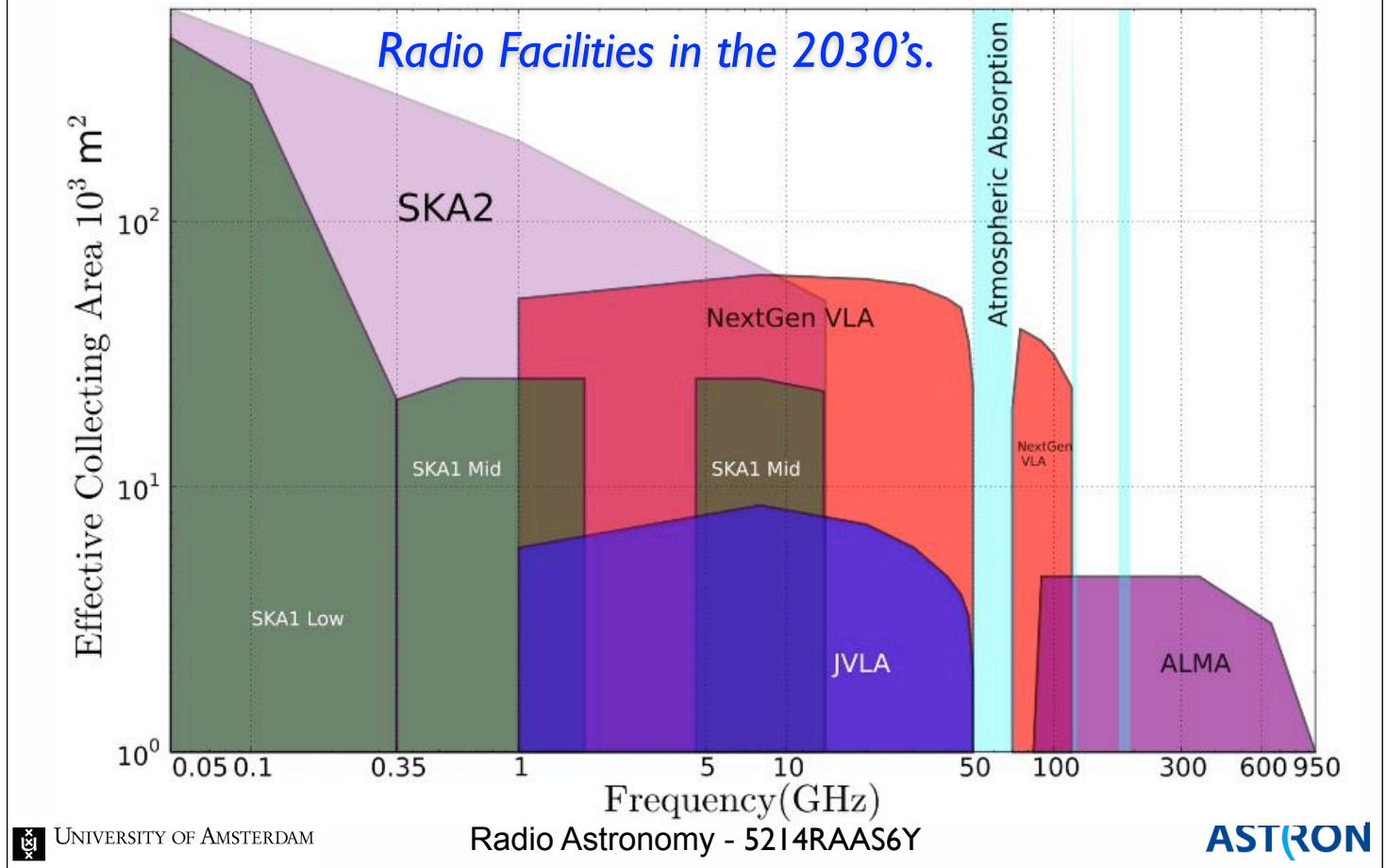
Sensitivity Comparison



Comparison of the expected sensitivity for the SKA1 and SKA2 with JVLA and LOFAR.

The SKA will be orders of magnitude more sensitive than existing telescopes!

Sensitivity Comparison

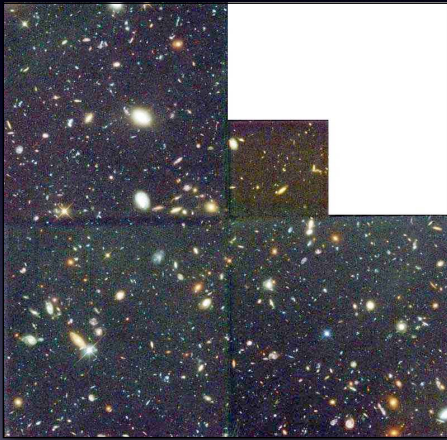


Radio Facilities in the 2030's.

Comparison of the expected sensitivity for the SKA1 and SKA2 with NGVLA and ALMA.

Expected Sensitivity

Hubble Deep Field



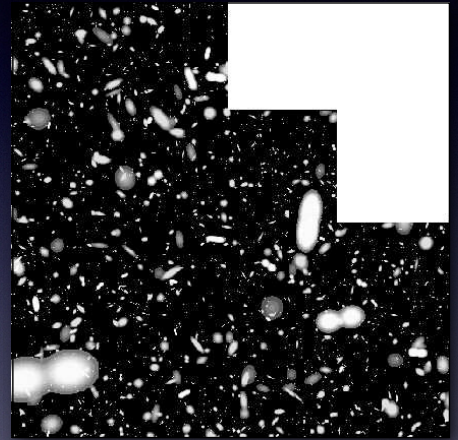
2.5 arcmin x 2.5 arcmin
~3000 galaxies

JVLA



50 hours at 8.7 GHz gives
6 sources at $>12 \mu\text{Jy}$

Simulated SKA



1000's of sources at $1 \mu\text{Jy}$ @
1.4 GHz (fraction of total FoV)

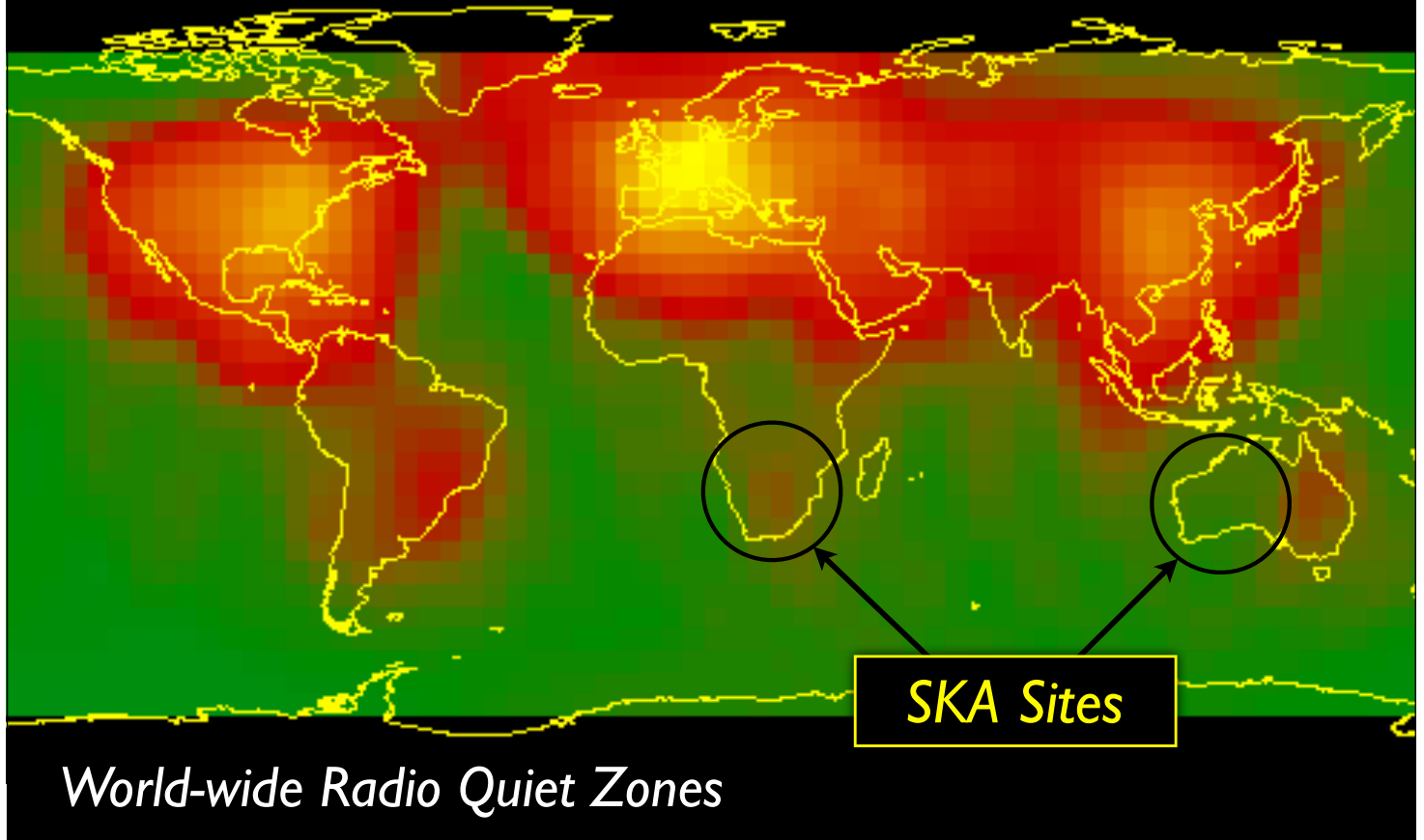


A comparison showing the difference in the number of galaxies seen by the VLA and the SKA.

A similar exposure with the SKA will yield 1000's of times more galaxies than current telescopes.

It will also have much bigger fields of view than the HST image shown and the VLA.

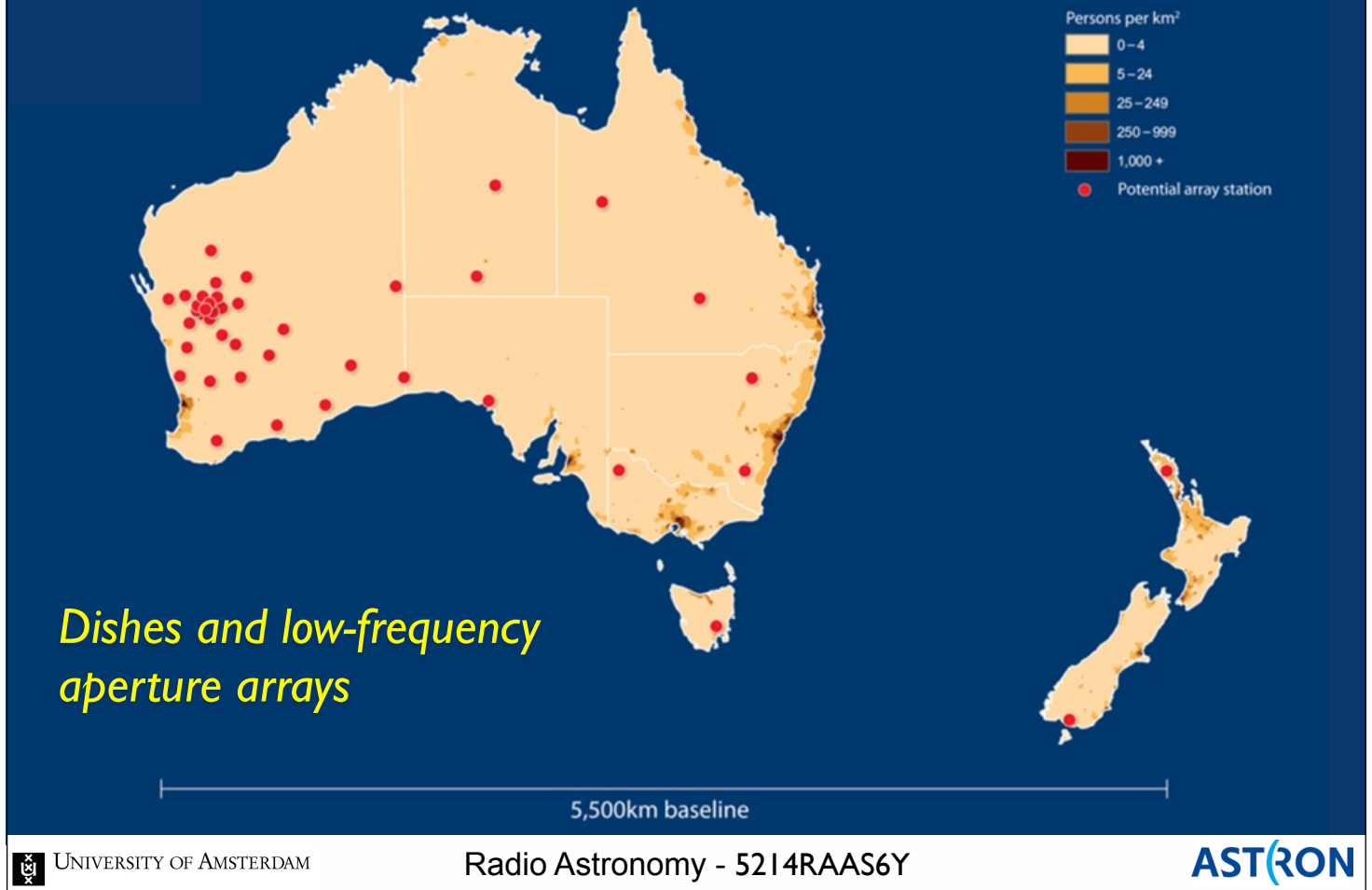
Where do we build it?



This image shows a map of the intensity of RFI around the world. Yellow areas have the strongest RFI.

Note the blazing yellow over the Netherlands...ouch. The deserts of South Africa and Western Australia however are the most radio quiet areas on earth.

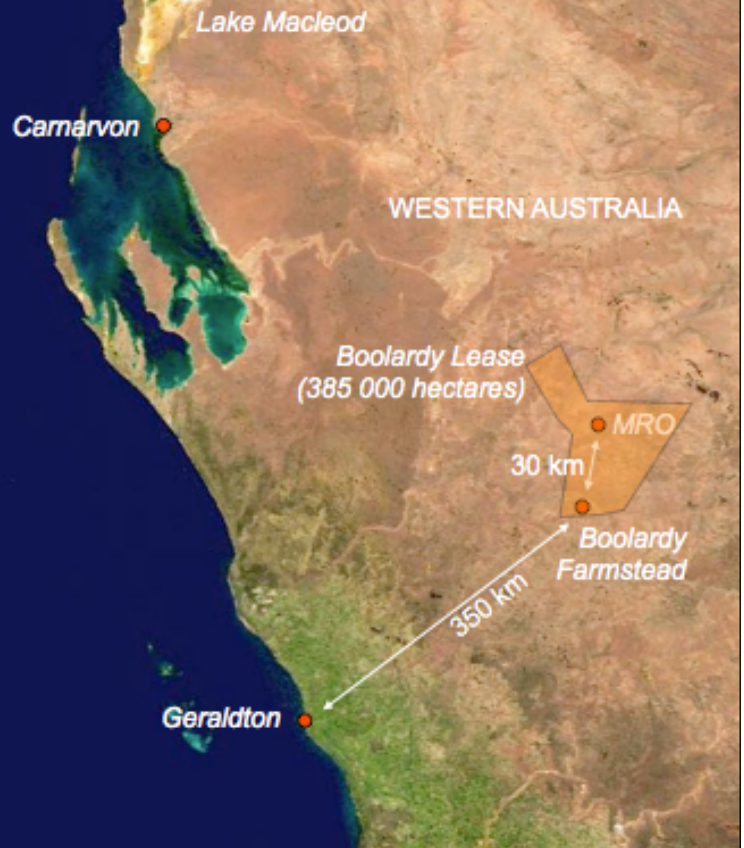
SKA Site I: Australia



The desert in western Australia will host the low-frequency part of the SKA. Originally it would have hosted a large array of dishes as well, but the design has changed recently to save money. The population density in this region is less than 4 people/km² or 0.4 nanohumans/cm².

Australian SKA Pathfinder (ASKAP)

- Shire of Murchison
- 50,000 km²
- 0 incorporated towns
- 29 sheep/cattle stations
- Population: 110



The low population density is great for RFI, it also makes every other aspect of building and running an observatory much more difficult.

No roads, water, or power, and many hours from any civilization.

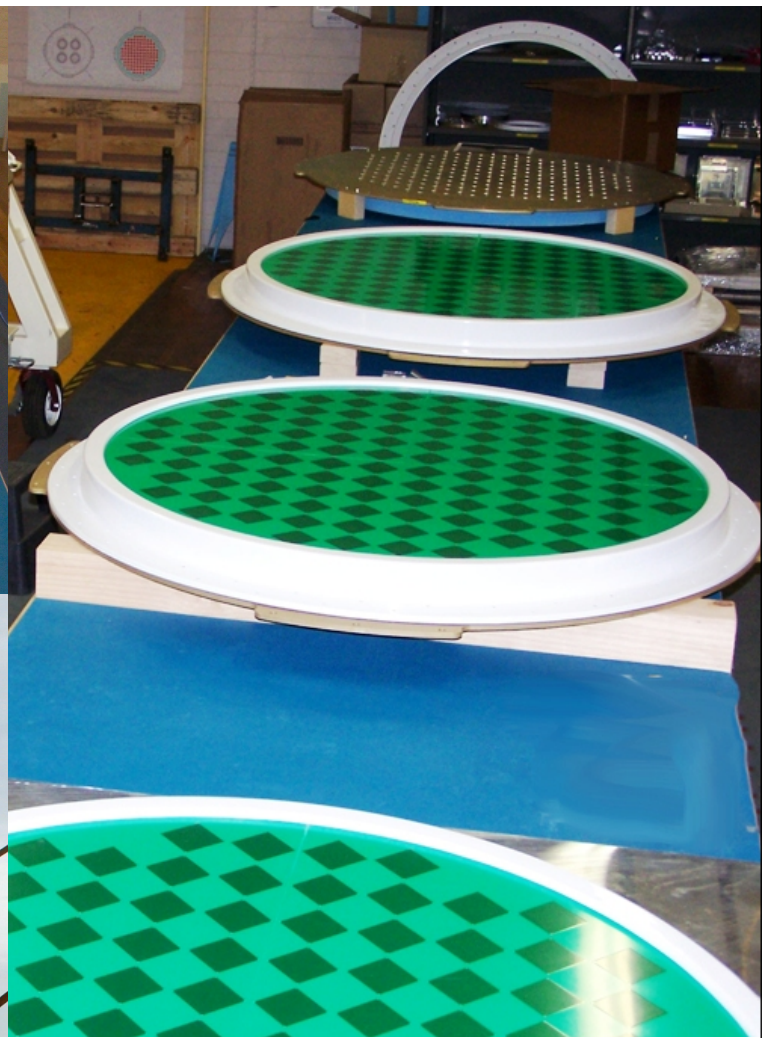
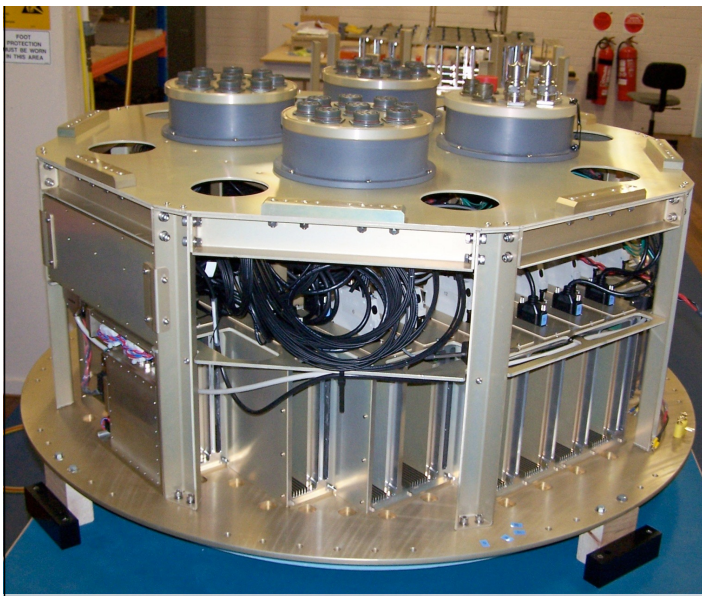
Constructing and maintaining the site infrastructure will be a big part of the cost of the SKA.



A fairly recent image of the Australian site.
The dishes you see are part of the Australian SKA
Pathfinder (ASKAP) array.



Close-up image of the ASKAP dishes. Notice the green circles in the focus. These are phased-array feeds (PAFs) similar to the APERTIF units being installed on WSRT. Like APERTIF, these PAFs will make ASKAP a powerful survey telescope especially for HI surveys.



Close-up image of the ASKAP PAFs.

Australian SKA Pathfinder

Design goals:

- **High-dynamic range imaging**
- **Wide field-of-view science**

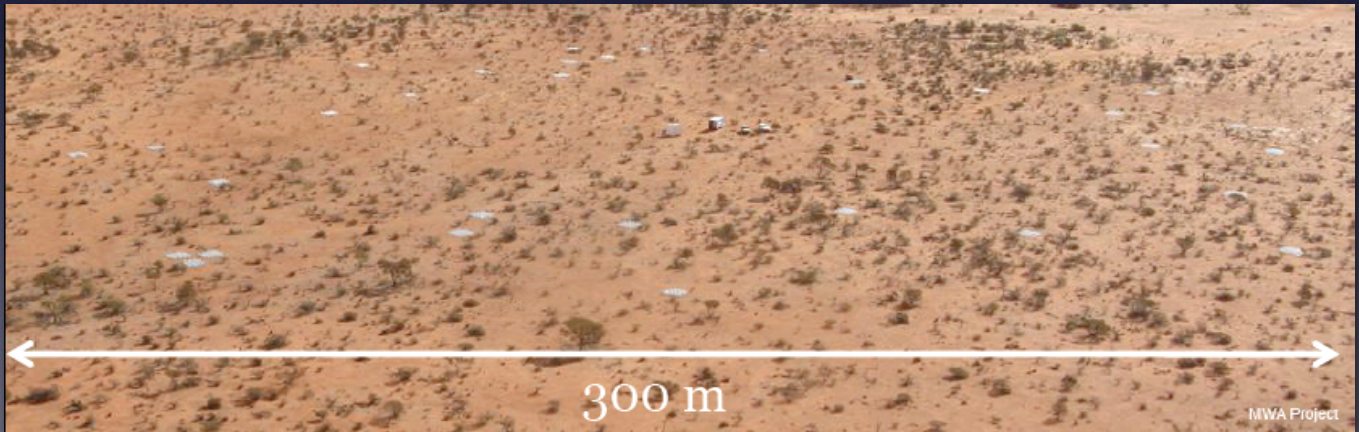
Number of dishes	36
Dish diameter	12 m
Maximum baseline	6 km
Resolution	30"
Sensitivity	65 m²/Kelvin
Survey Speed	1.3x10⁵ m⁴/kelvin²/deg²
Tsys/η	63 Kelvin (e.g. Tsys = 50K, η = 80%)
Observing frequency	700 – 1800 MHz
Field of view	30 deg²
Processed bandwidth	300 MHz
Spectral channels	16384
Focal Plane Phased Array	188 channels (94 beams)

System characteristics for the ASKAP array.
Notice the relatively short baseline length.
ASKAP will do relatively deep HI surveys, but at relatively modest resolution.

Murchison Widefield Array (MWA)

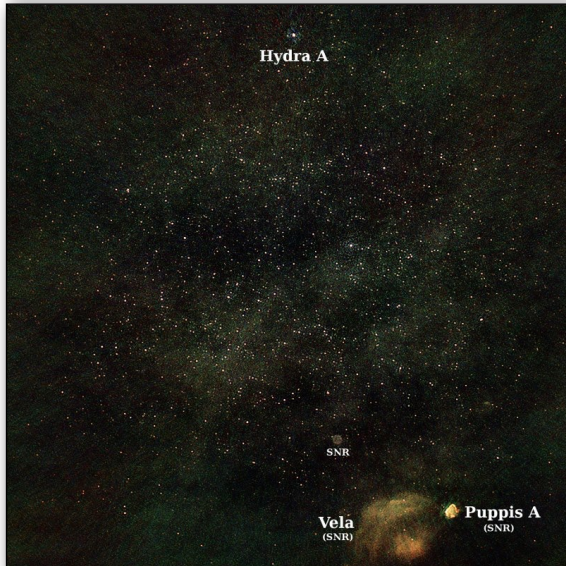
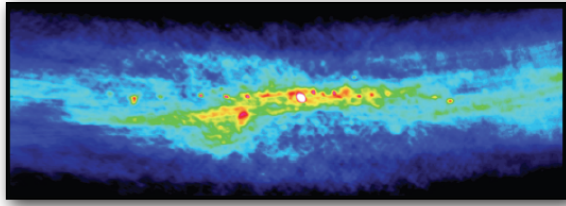


- Low-frequency AA
- 128 tile array
- SKA low precursor



In addition to the ASKAP mid-frequency SKA pathfinder telescope, there is a low-frequency pathfinder being developed in Australia as well. The MWA is similar to the technology used in LOFAR and covers a similar piece of the low-frequency spectrum. It lacks the longer baselines and does not go quite as low in frequency. The combination of LOFAR and MWA will give us a relatively deep survey of the entire sky.

Murchison Widefield Array (MWA)



Parameter	Value
Frequency range	80 – 300 MHz
Number of receptors	2,048 dual-pol. dipoles
Number of “tiles”	128 (expandable to 256)
Number of baselines	8,128
Collecting area	~2500 m ² @ 200 MHz
T _{sys}	25 K (Rx); 125 K (sky) @ 200 MHz
Field of view (diameter)	~40° @ 200 MHz
Configuration	Core: 1.5 km in diameter (87% of collecting area) Extended: 3 km in diameter (13% of collecting area)
Bandwidth	220 MHz sampled, 30.72 MHz processed
Spectral channels (correlator)	768 (40 kHz resolution)
Temporal resolution (correlator)	0.5 s uncalibrated; 8 s calibrated
Polarisations correlated	Full Stokes
Continuum point source sensitivity	80 mJy in 1 s @ 200 MHz 1.3 mJy in 1 hr @ 200 MHz
Voltage capture for full array	yes

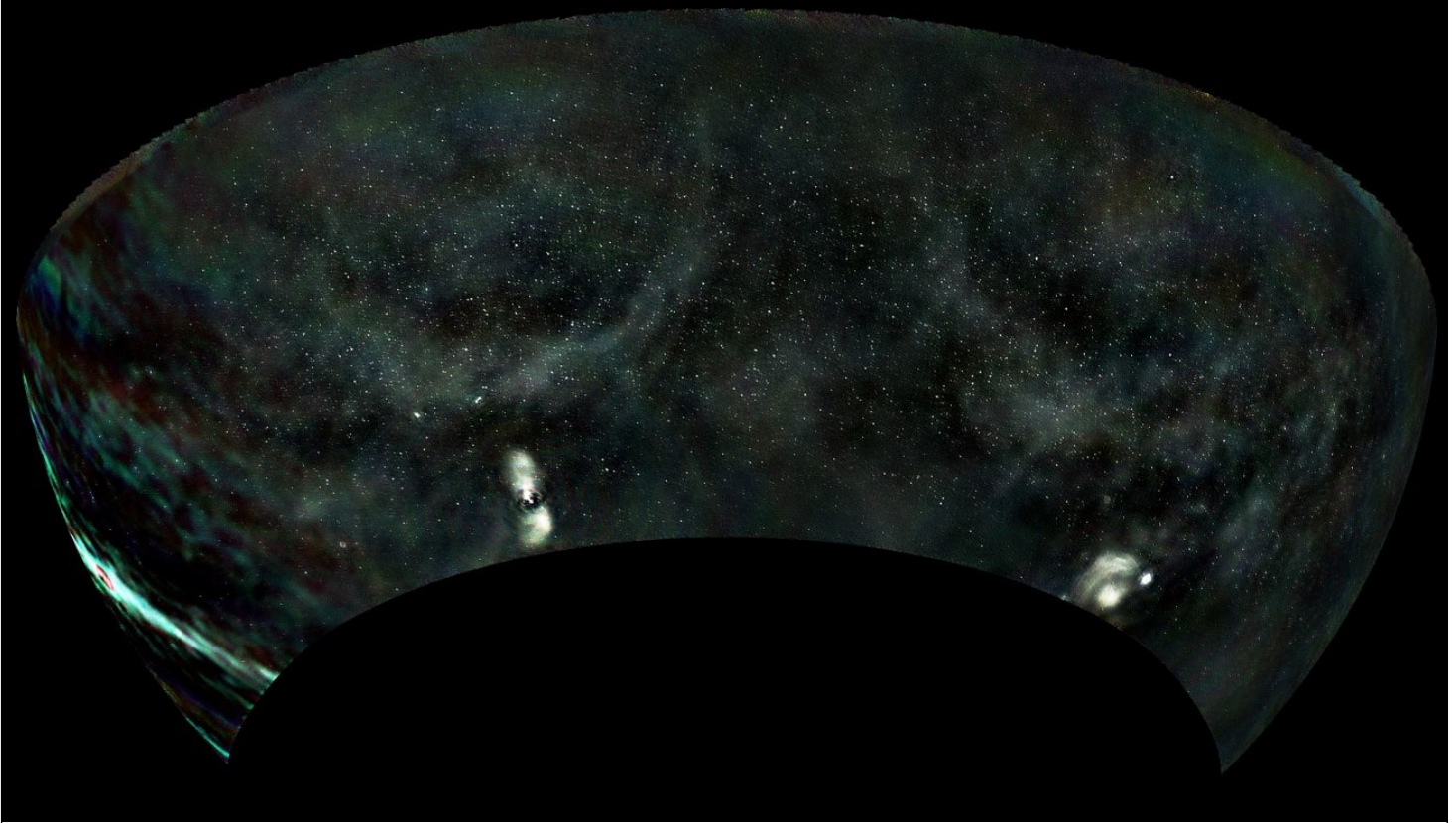
System characteristics for the MWA.

Notice that it does not go as low in frequency as LOFAR and the much baselines.

MWA was originally intended as a specialized EoR experiment.

The short baselines and good instantaneous uv coverage are well suited for EoR observations.

Murchison Widefield Array (MWA)



Example of the an all-sky map of the southern sky from MWA.

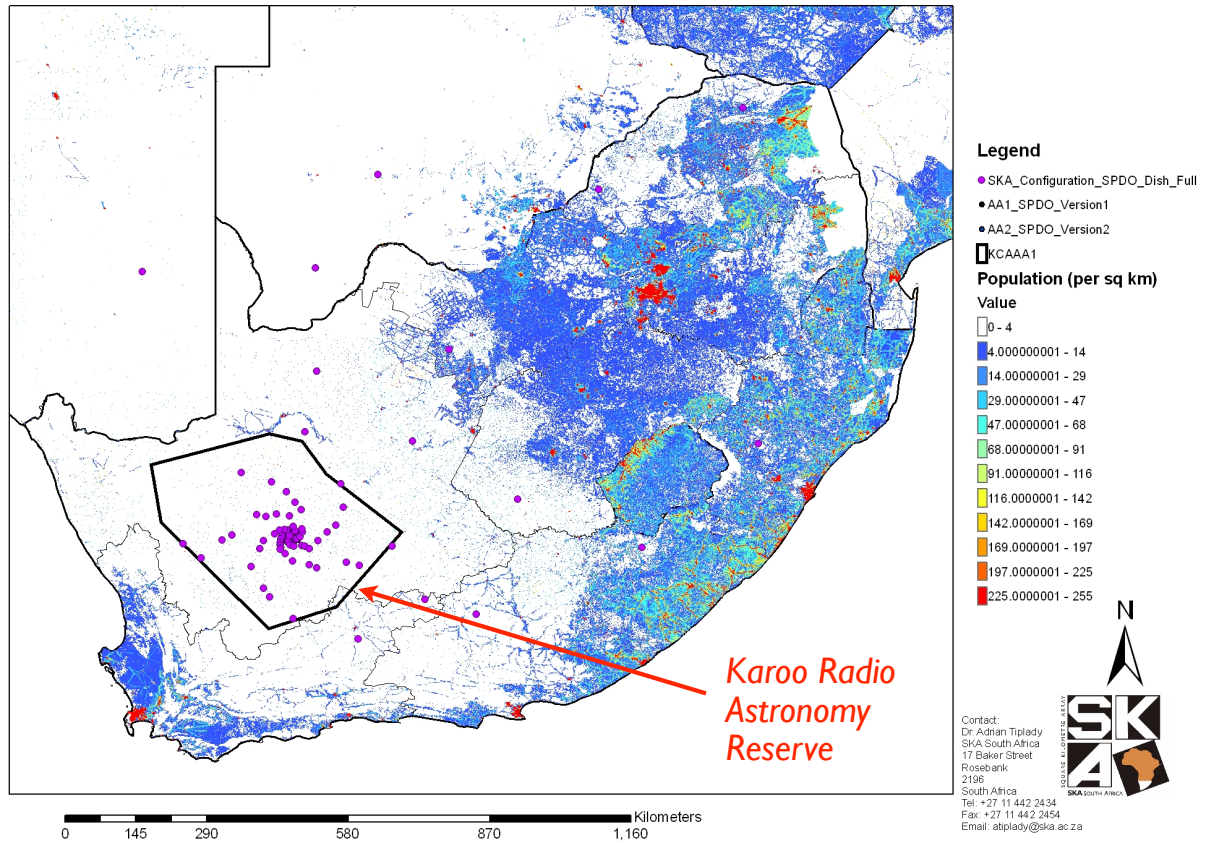
SKA Site 2: South Africa

*Dishes and
mid-frequency
aperture arrays*



The second, higher frequency component of the SKA will be based in South Africa.

SKA Site 2: South Africa



The desert in South Africa has similar advantageous properties to Western Australia.
Very low population density and the very low RFI that goes with it.
This remote site has all the same logistical difficulties as well.



A fairly recent image of the South Africa desert.

MeerKAT SKA Pathfinder



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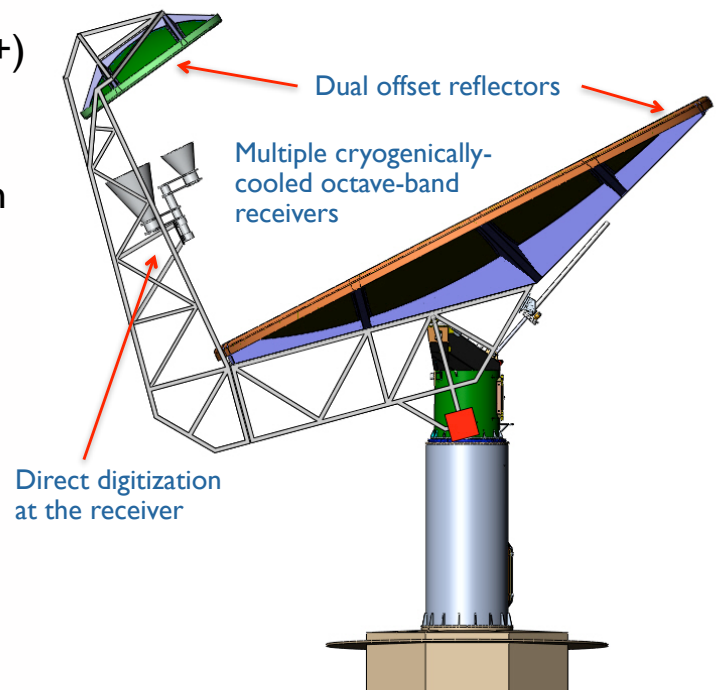
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A fairly recent image of the South Africa site.
The dishes you see are part of the South African SKA
Pathfinder MeerKAT array.

MeerKAT SKA Pathfinder

- Will be most sensitive cm-wavelength telescope in the southern hemisphere
- 580 MHz – 15(+) GHz (i.e. SKA-mid +)
- Imaging and non-imaging modes
- High filling factor for baselines < 1 km
- 64 x 13.5 m gregorian offset antennas
- 8 km maximum baseline
- 70 % in < 1km diameter core
- Future expansion to 20+ km
- KAT7 prototype array in place

SKA Baseline Design



System characteristics for the MeerKAT array.
Like ASKAP, notice the relatively short baseline length.

MeerKAT will do relatively deep HI surveys, but at relatively modest resolution.

MeerKAT SKA Pathfinder



Computer rendered image of the anticipated SKA dish design.

PAPER (Precision Array to Probe Epoch of Reionization)

Karoo Desert



- Low-frequency AA
- 64 tile array
- EoR experiment

NRAO Greenbank



National Radio Astronomy Observatory
A facility of the National Science Foundation



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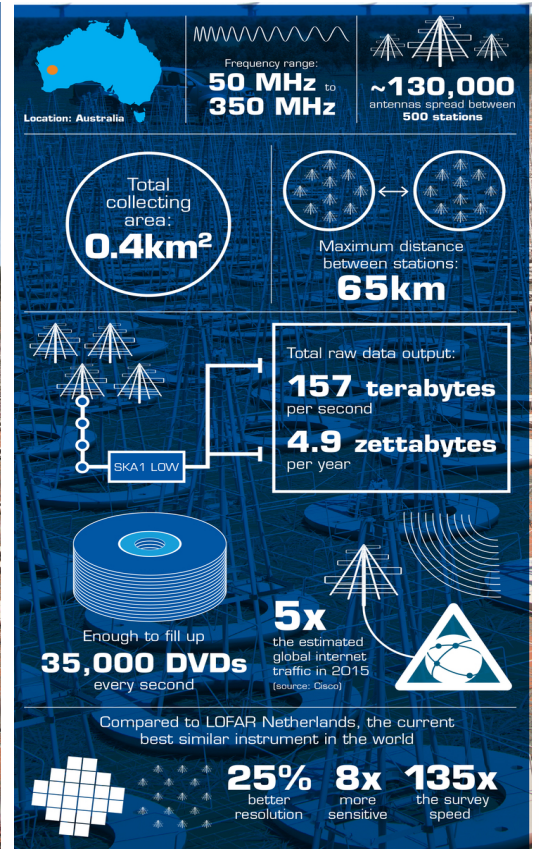
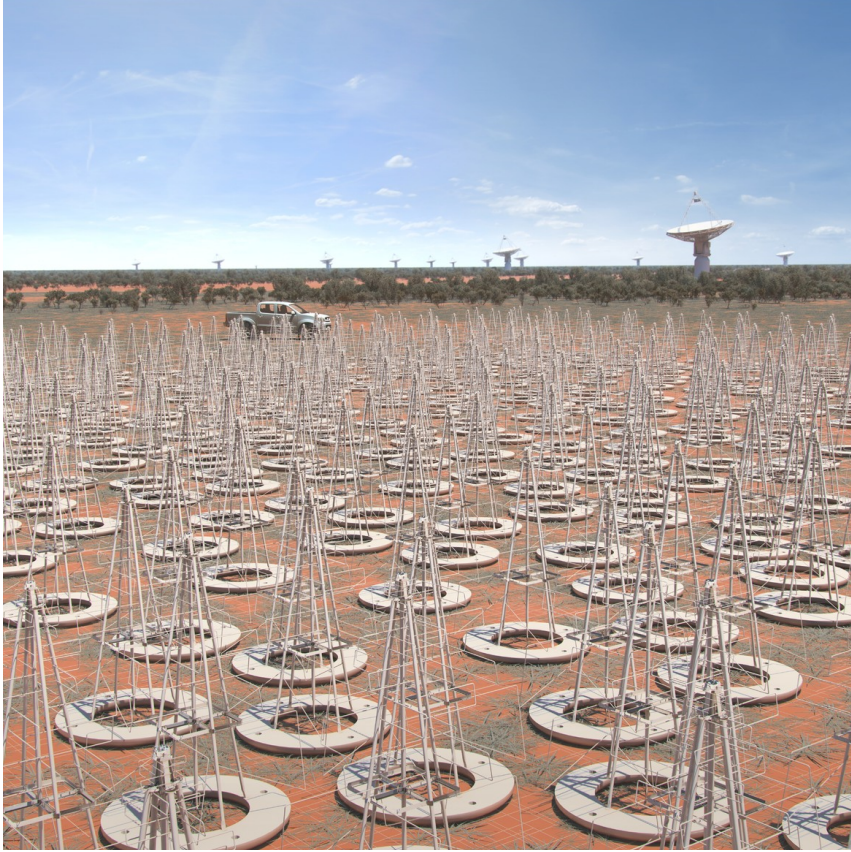
There is also a low-frequency phased array prototype at the SA SKA site.
PAPER is a specially designed experiment to search for the EoR signal.

Computational Challenges

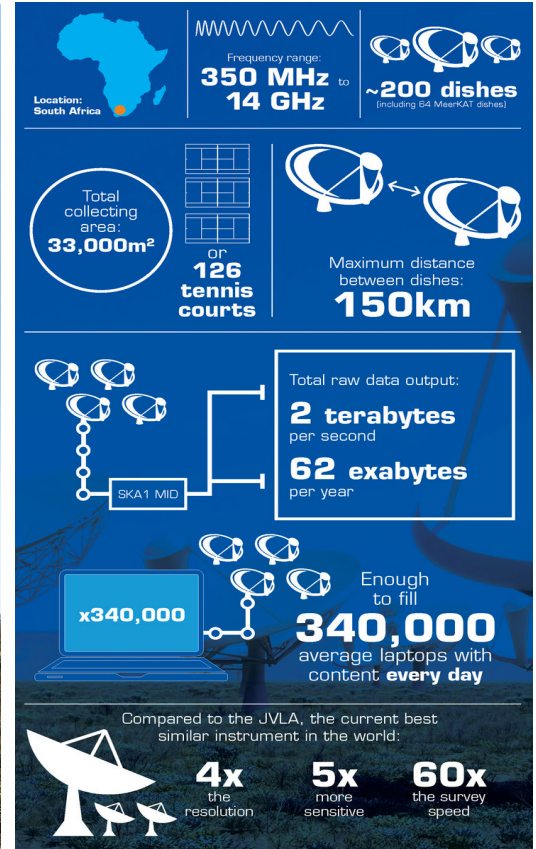
The new technologies that make telescopes like LOFAR and ultimately the SKA possible require heavy computation.

Computation in this sense means storage of lots of data, lots of processing to do something useful with this data, and networking to transport that data. These challenges are some of the biggest issues facing the success of the SKA and the future of radio astronomy.

SKA-LOW Data Rates



SKA-Mid Data Rates



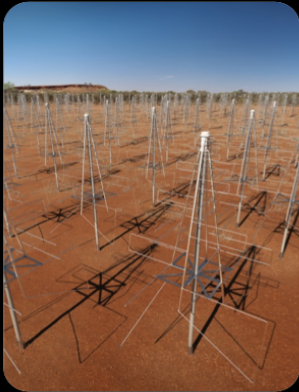
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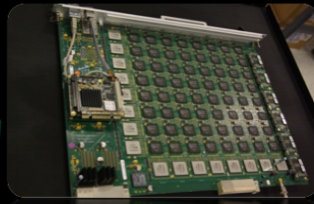
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SKA System Data Flow

Antennas



Digital Signal Processing (DSP)



Transfer antennas to DSP
2020: 5,000 PBytes/day
2030: 100,000 PBytes/day

Over 10's to 1000's kms

To Process in HPC
2020: 50 PBytes/day
2030: 10,000 PBytes/day

Over 10's to 1000's kms



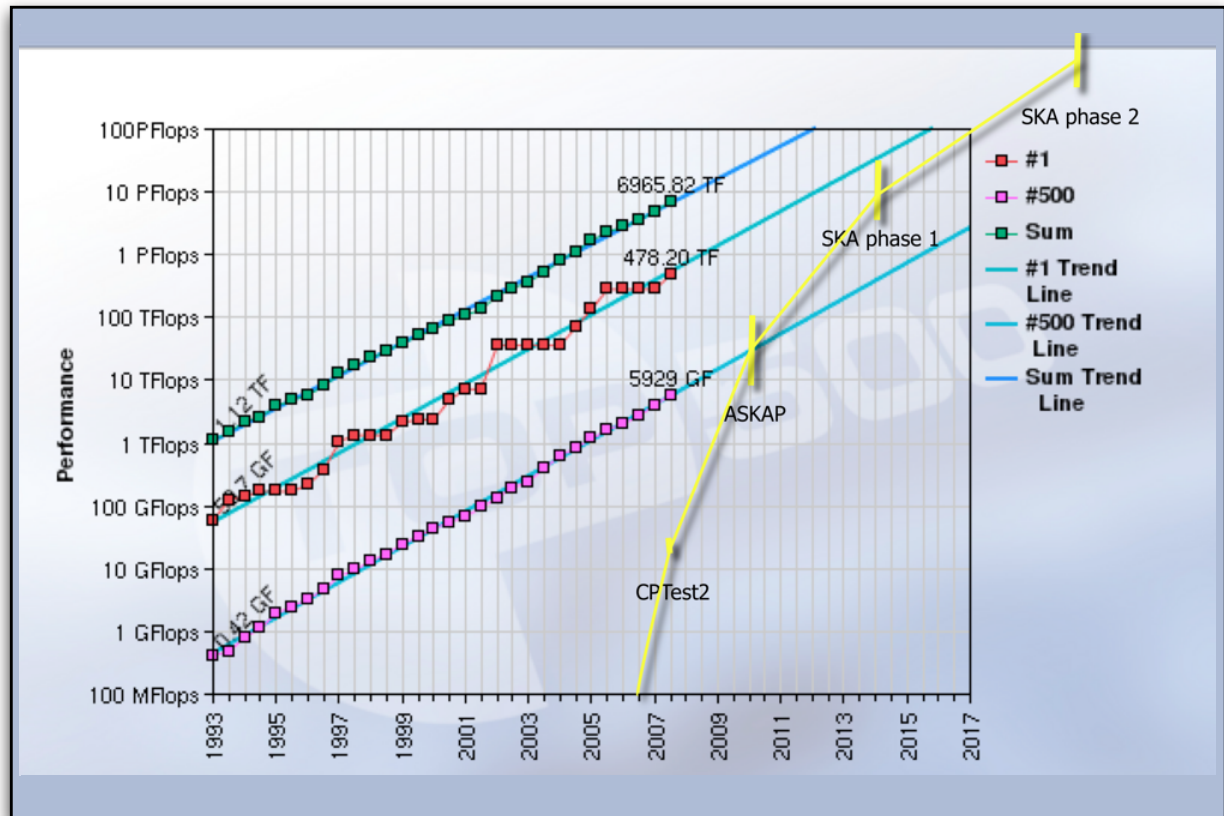
High Performance Computing Facility (HPC)

HPC Processing
2020: 300 PFlop
2028: 30 EFlop

30 – 300 Pbytes / year of fully processed science data products!



SKA Computing Costs



The growth of computing power over the last few decades. The curves show the fastest and fastest 500 computer systems as a function of time. The yellow line shows the estimated total computing power needed for the SKA.

SKA Data Products

Experiment	T_{obs}	B/km	D/m	N_b	N_{ch}	N_v	Size / TB
High resolution spectral line	3600	200	15	1	32000	$5 \cdot 10^{13}$	200
Survey spectral line medium resolution	3600	30	56	1000	32000	$8 \cdot 10^{13}$	330
Snapshot continuum – some spectral information	60	180	56	1200	32	$7 \cdot 10^{12}$	30
High resolution long baseline	3600	3000	60	1	4	$7 \cdot 10^{14}$	360

- **~0.5 – 10 PB/day of image data**
- **Source count ~ 10^6 sources per square degree**
- **~ 10^{10} sources in the accessible SKA sky, 10^4 numbers/record**
- **~1 PB for the catalogued data**

100 Pbytes – 3 EBytes / year of fully processed data

This slide shows some typical expected data volumes for a few typical SKA observations.

A *single* observation could easily generate 100's of TB of data.

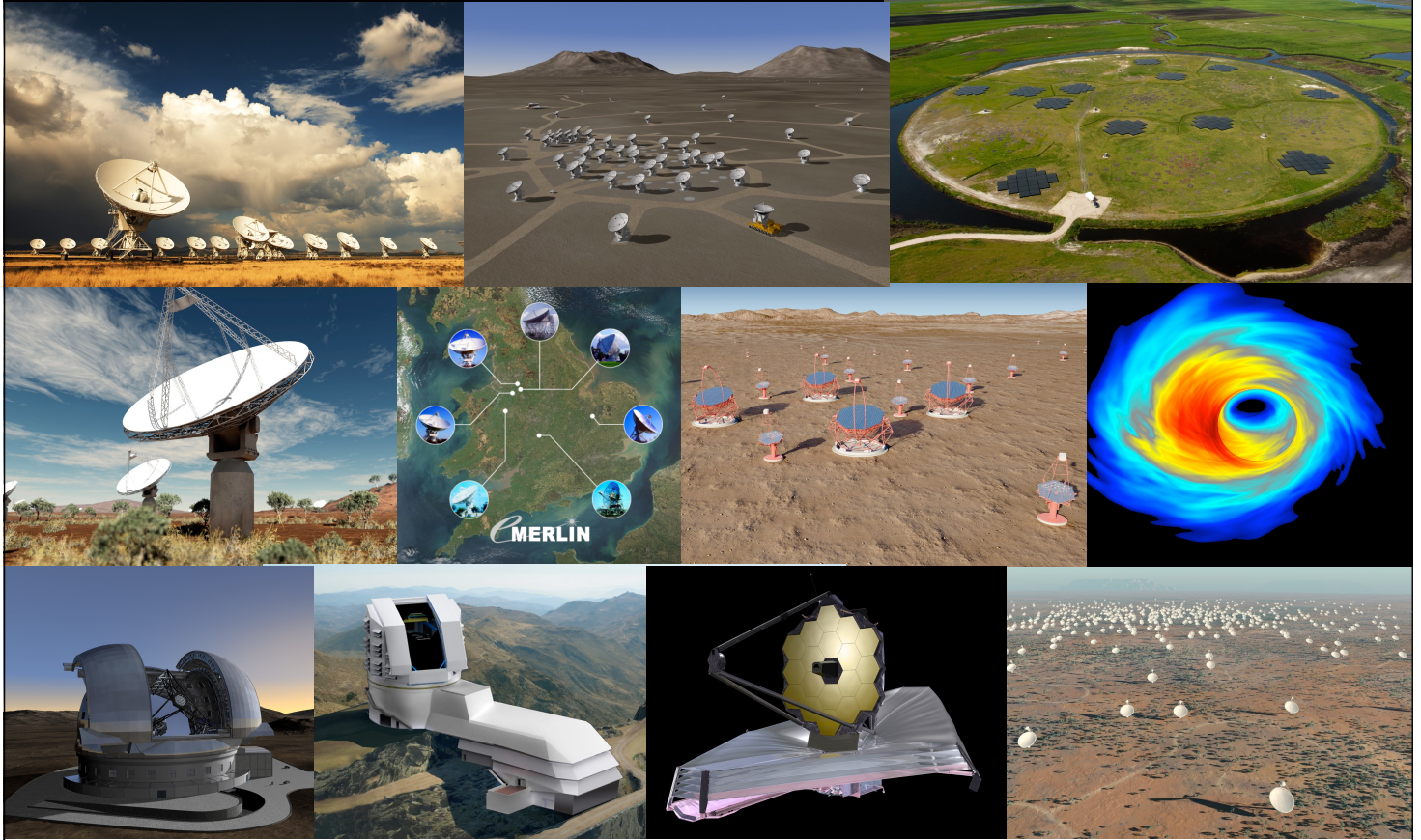
At this level, the entire paradigm for how we do astronomy changes.

No more loading data onto a laptop and reducing it at home.

Intermission

Data Intensive Astronomy

Data Intensive Astronomy



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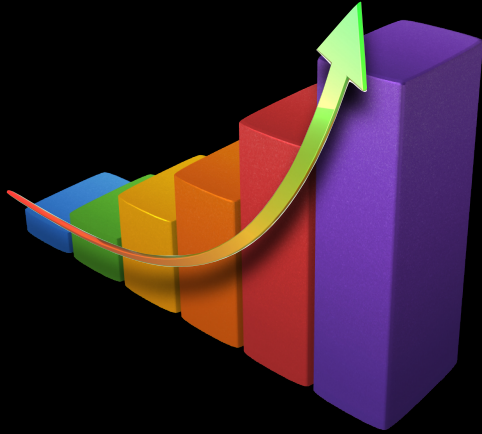
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Data intensive astronomy is not just an issue for radio astronomy.
Many new telescopes at a variety of wavelengths will generate huge datasets.
Simulated datasets can also be very large and create the same sorts of issues.

Data Intensive Astronomy

Exponential
Growth of
Data Volumes



...and
Complexity

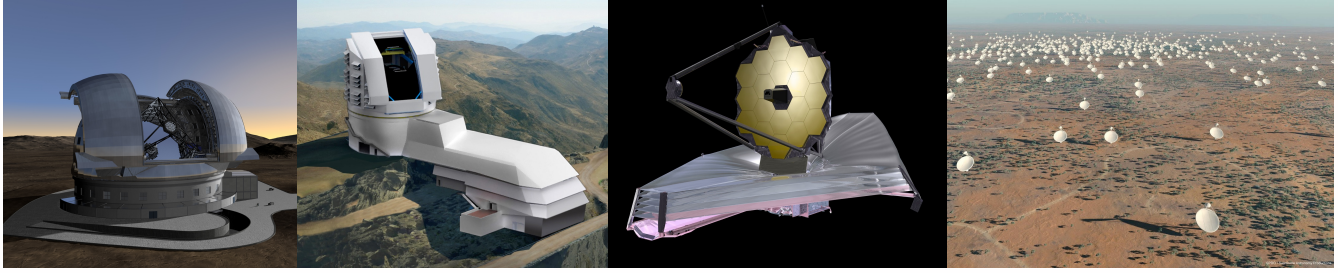


Understanding of
complex phenomena
requires complex data

- From data poverty to data glut
- From data sets to data streams
- From static to dynamic, evolving data
- From centralized to distributed resources

Data intensive astronomy is not just about computing, storage, and networking. It is also about the nature of the data itself which has also evolved. Its about data volume but also the complexity of the data itself.

What does “Data Intensive” mean?



- Science is increasingly driven by large data sets
- Data collection in large collaborations
- Analysis done on the archived data
- New instruments will produce petascale datasets

Petascale analysis require exascale data management!

The Era of Big Surveys

*Modern sky surveys obtain $\sim 10^{12} - 10^{15}$ bytes of images
Catalogs $\sim 10^8 - 10^9$ objects (stars, galaxies, etc.)
and measure $\sim 10^2 - 10^4$ numbers per object*

We've repeatedly discussed how the next generation of radio telescopes have been optimized to perform surveys.

This trend toward all-sky surveys is also going on at other wavelengths especially in the optical.

Even the final products of these surveys, i.e. images and source catalogs, can be very large.

LOFAR Surveys in Context

2000.- 2014 Sloan Digital Sky Survey (SDSS)

120 Mpixel camera, (0.08 PB in 10 yrs)

3×10^8 unique sources (4 TB)

2018 - 2028 Large Synoptic Survey Telescope (LSST)

3.2 Gpixel camera (6 PB per year)

1000's observations of every source

few $\times 10^9$ sources, few $\times 10^{12}$ rows (10 PB)

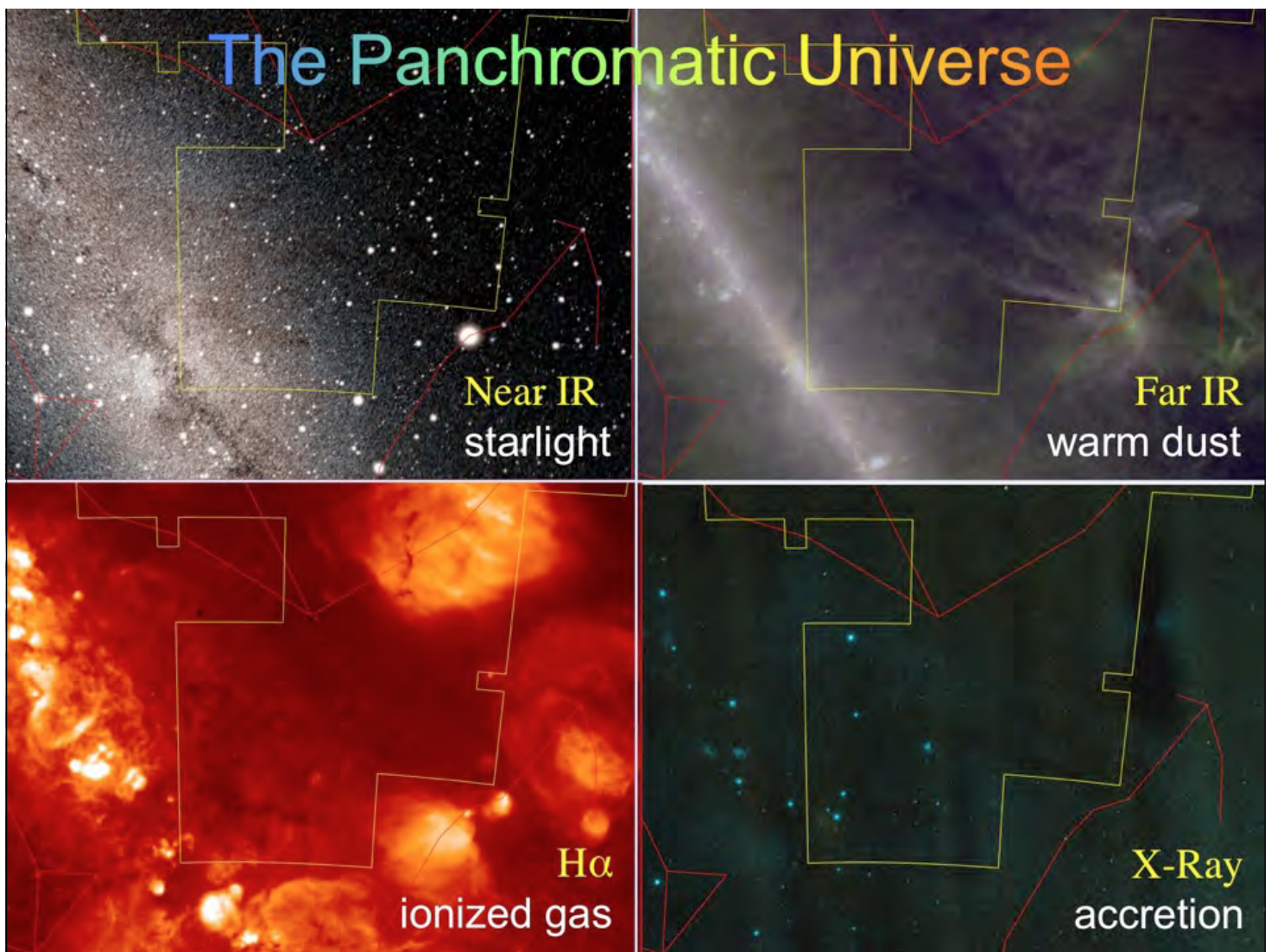
2013 - 2018 LOFAR Low-Frequency Sky Survey (LFSS)

$\sim 100 \text{ deg}^2$ FOV ($\sim 5.2 \text{ Gpixel}$) ($\sim 0.1 \text{ EB}$ per year)

$10\text{-}10^3$ freqs, $10^2\text{-}10^4$ observations of every source

few $\times 10^8$ sources, few $\times 10^{12}$ rows ($\sim 1\text{-}5 \text{ PB}$)

A comparison of the size of next generation surveys in the optical compared to LOFAR.
Even the source catalogs can reach sizes of petabytes.



For many types of science, it is important to combine data across many wavelengths.

To be able to make these sort of multi-wavelength analysis, the data must be archived and we need a way to easily combine them. At the moment, we do these sorts of analyses manually.

At the scale of the SKA and LSST, we won't be able to do it manually any longer.

Numerical Simulations

Complex theoretical models are expressed as numerical simulations

...and then must be matched against large collections of observational data

Bolshoi Simulation (Klypin et al. 2011)

⇒ **Can easily generate TBs per hour**

Numerical simulations are a new kind observational data.

They can be as equally large and complex as data obtained from telescopes.

By making it accessible through archives, it can be compared to actual observations and used in analysis.

Data Intensive Radio Astronomy

- Radio astronomy is already data intensive
- Current facilities already generating large data (JVLA, ALMA, eMERLIN, LOFAR, etc.)
- New facilities orders of magnitude larger (MeerKAT, ASKAP, and SKA)



LOFAR

Long Term Archive

28PB



6PB

United States
Census
Bureau

4PB

NASDAQ

3PB

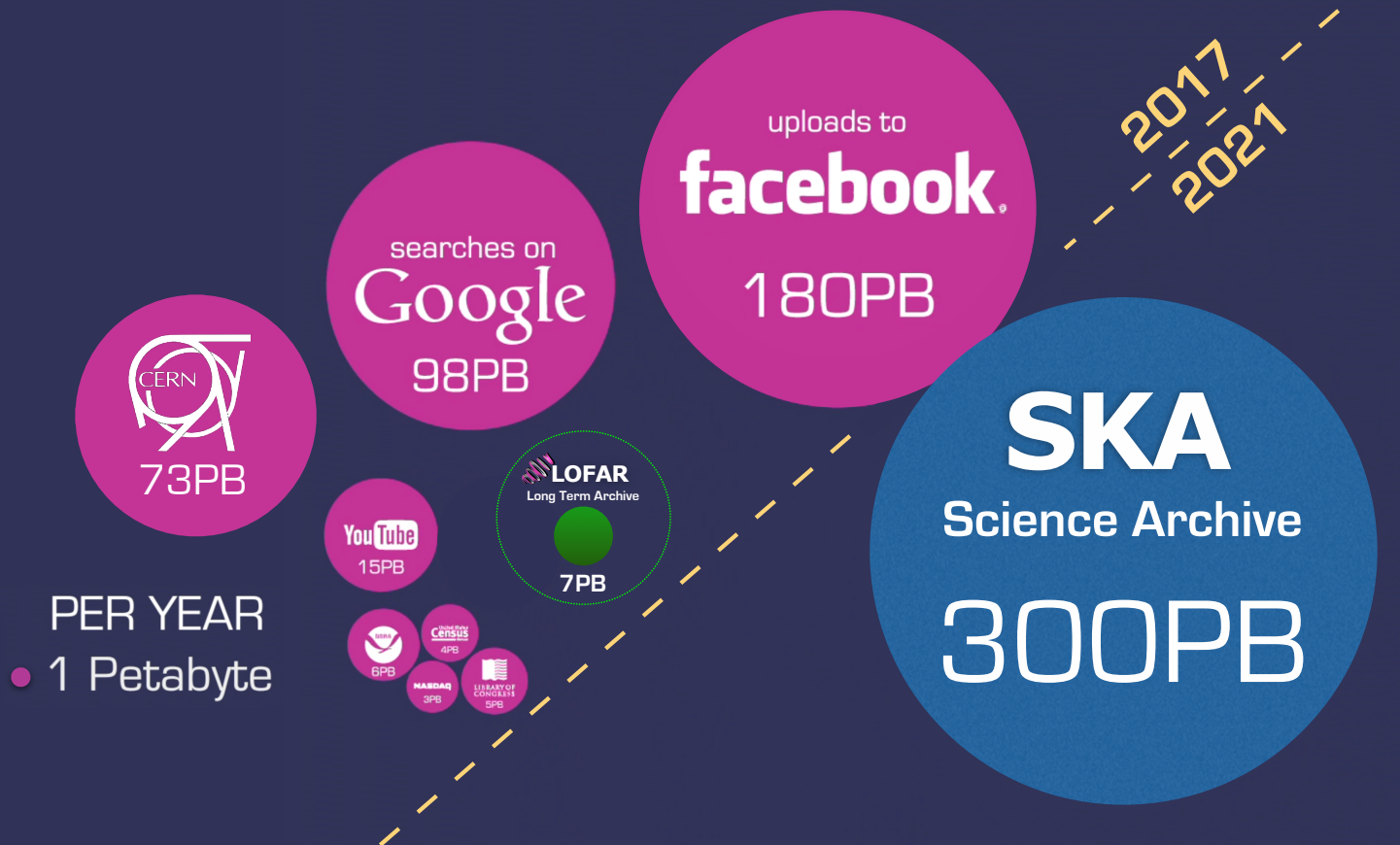


LIBRARY OF
CONGRESS

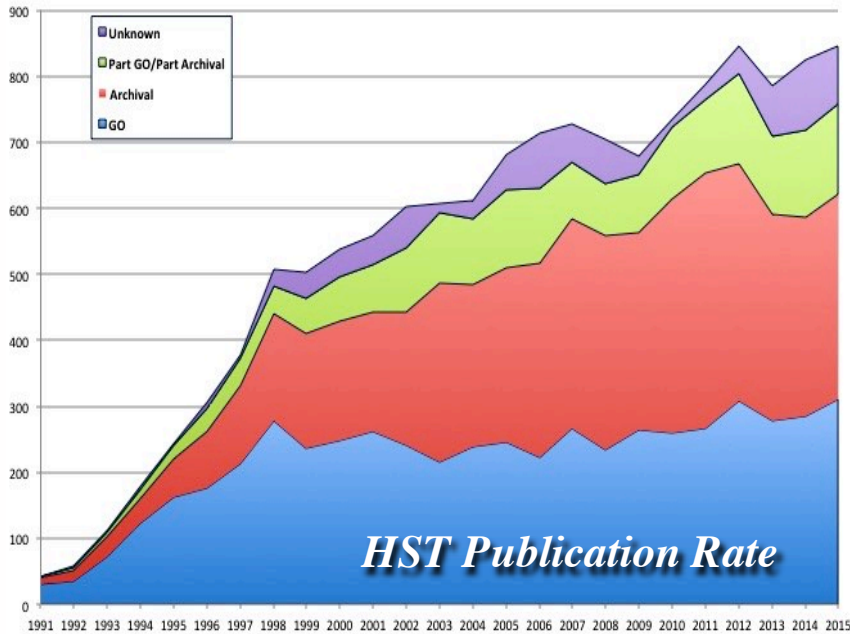
5PB

1 Petabyte [PB]

Future SKA Science Archive



Impact of Science Archives

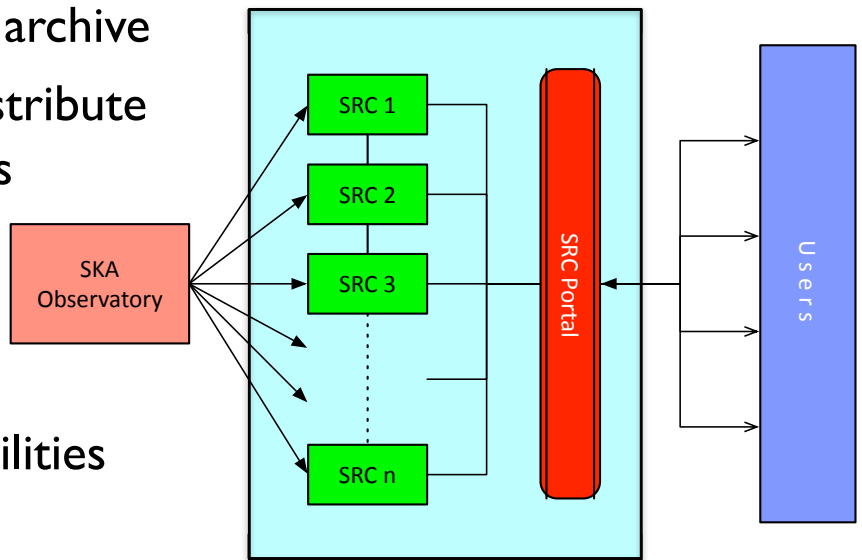


- Assumes the archives are persistent and maintained
- Assumes archival data is open and accessible to users
- Assumes data products stored are appropriate for general use
- Assumes users retrieving data have resources to process to a science result

Archives can increase the total scientific output from a telescope!

SKA Regional Centres

- SKA Regional Centres (SRCs) will host SKA science archive
- Provide access and distribute data products to users
- Provide access to compute and storage resources
- Provide analysis capabilities
- Provide user support
- Multiple regional SRCs acting as a single network

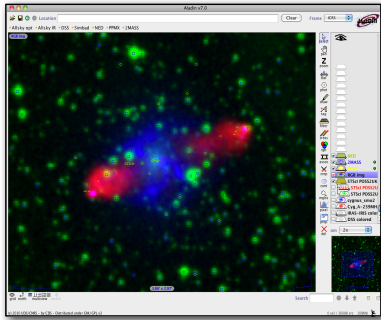


Primary interface for SKA data analysis

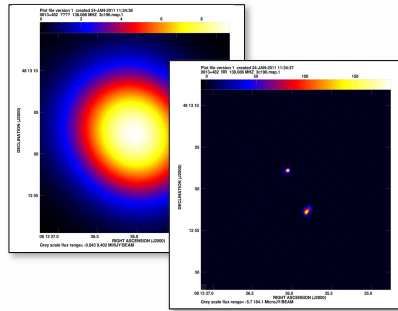
Data Centre Functionality

Data Discovery

- Observation database
- Associated metadata
- Quick-look data products
- Flexible catalog queries
- Integration with VO



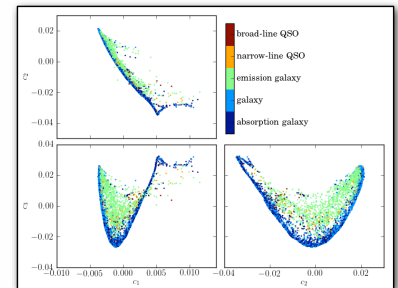
Data Processing



- Calibration and Imaging
- Mosaicing
- Source extraction
- Catalog creation
- DM searches

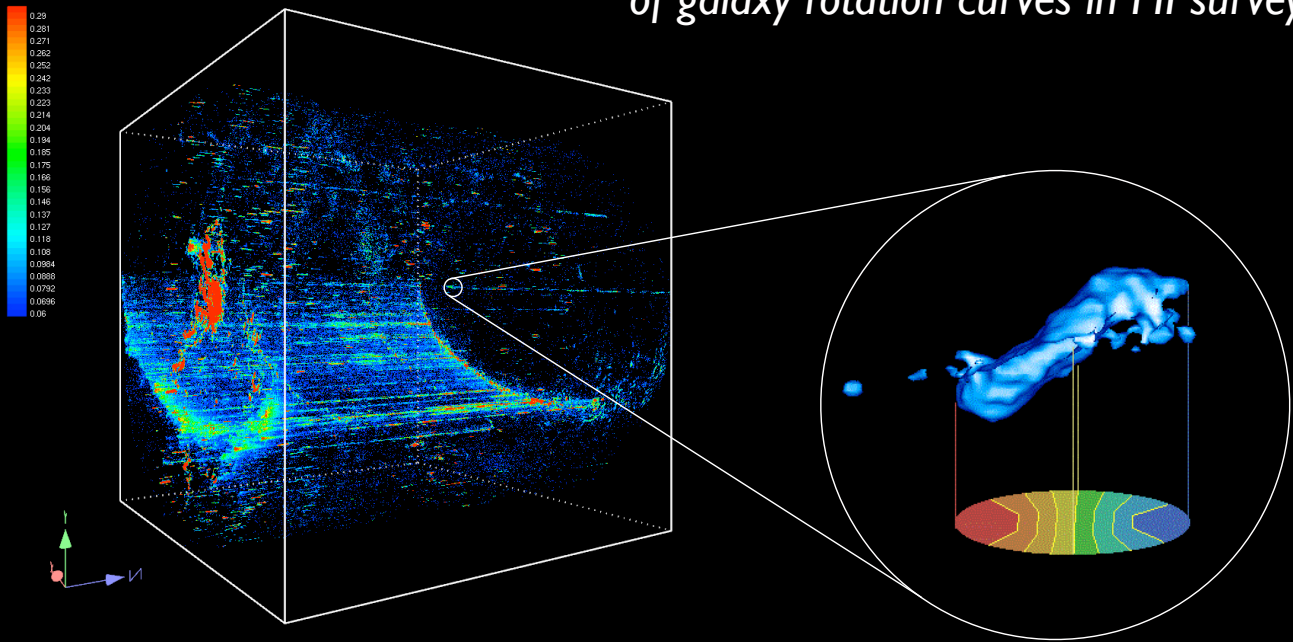
Data Analysis

- Multi-wavelength studies
- Catalog cross-matching
- Light-curve analysis
- Transient classification
- Feature detection
- Visualization



Visualization, Classification, Inference

*Automated detection and calculation
of galaxy rotation curves in HI surveys*



HIPASS HI Cube (Jurek et al. 2010)

SKA1-Mid cube ~ 0.9 PB

Global Network of Centres



Beyond the SKA



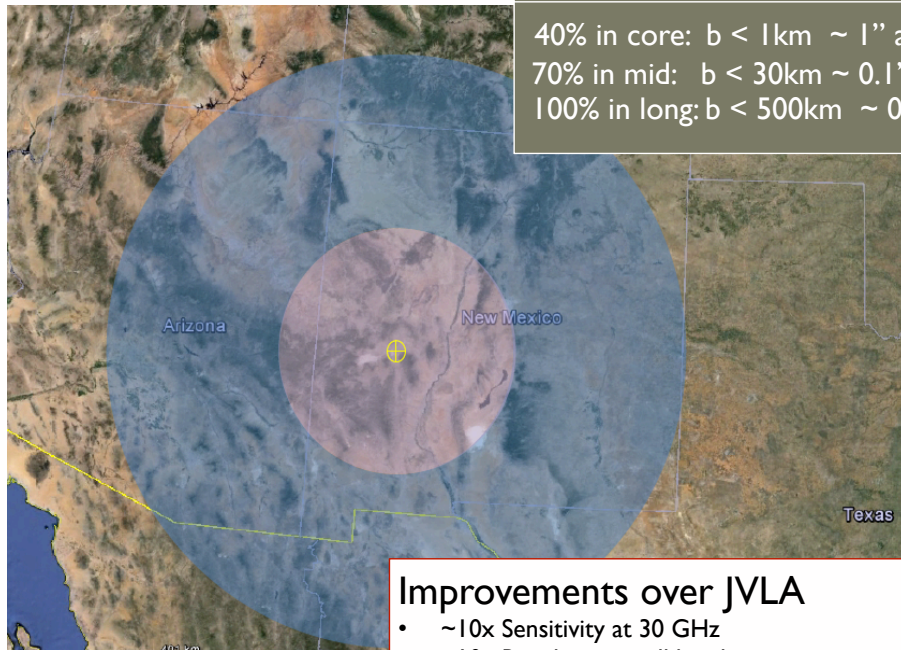
Next Generation VLA

**Thermal imaging at milli-arcsecond scale
resolution
10x effective collecting area & resolution of JVLA**

Next Generation VLA

JVLA: Good 3mm site, elev. $\sim 2200\text{m}$

40% in core: $b < 1\text{km} \sim 1''$ at 30GHz
70% in mid: $b < 30\text{km} \sim 0.1''$
100% in long: $b < 500\text{km} \sim 0.01''$

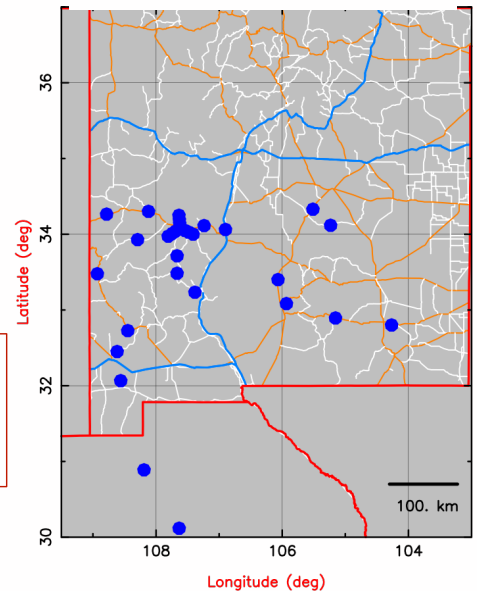


Improvements over JVLA

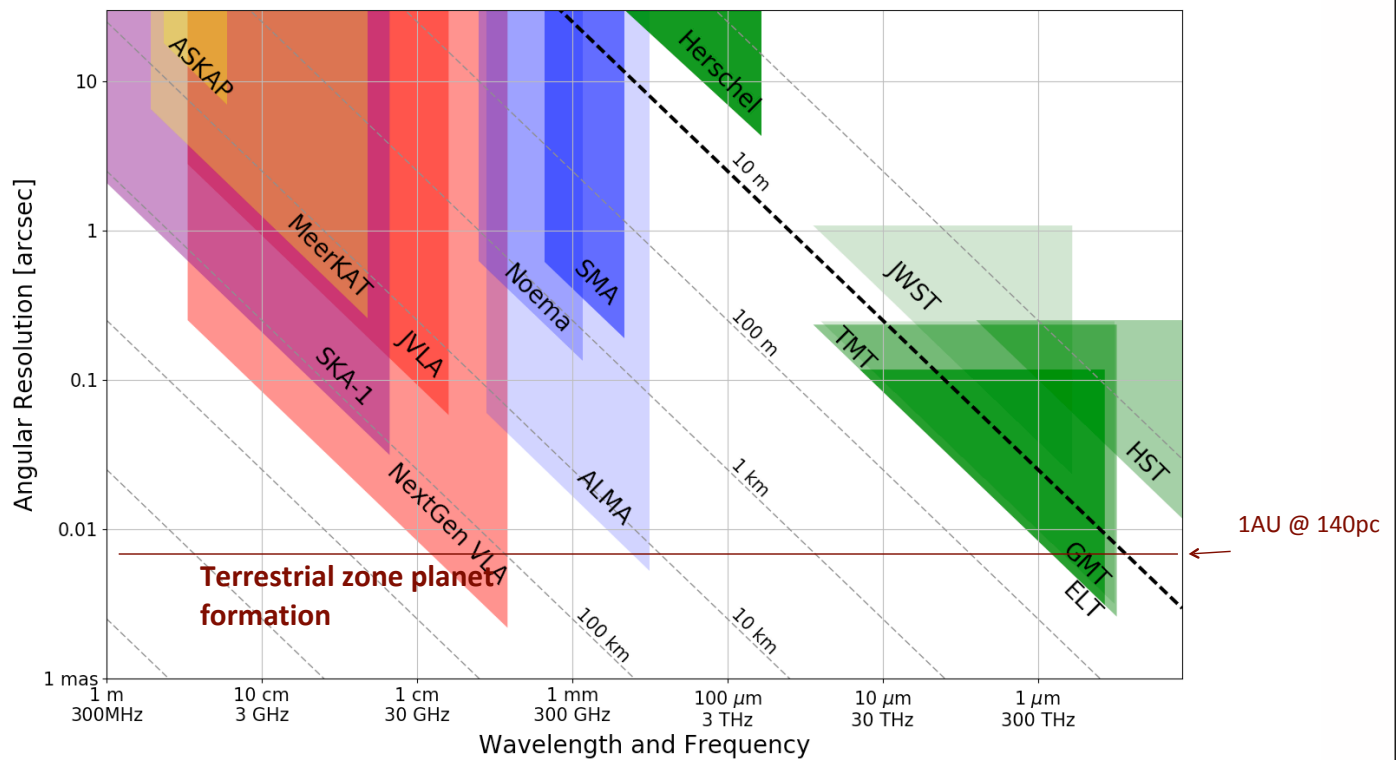
- $\sim 10\times$ Sensitivity at 30 GHz
- $\sim 10\times$ Resolution at all bands
- 1.2 – 116GHz (vs. 1 – 50GHz)

Located in U.S. and Mexico

Potential Station Locations



Bridging SKA & ALMA Scientifically



Begin construction in 2025, complete in 2034

Upgraded SKA

Sensitivity
RFI
Resolution
Computing

Add collecting area / bandwidth
Spectrum management / Removal
Size of Earth / Maximum frequency
Moore's law / Quantum computing

Extension of modern VLBI networks, leverages existing infrastructures

Where do we go next once the SKA has been built?
It would be fairly straight-forward to extend the existing technology to blanket the globe.
We could essentially build a radio telescope the size of the earth.
We would however face the same digital challenges that we are currently facing with the SKA.

Low-frequency Radio from Space

Netherlands – China Long-wavelength Explorer (NCLE)

**NL developed LF antenna
integrated on Chinese orbiter
to explore below 10 MHz**

Chang'e 4 Orbiter

Chang'e 4 Lander

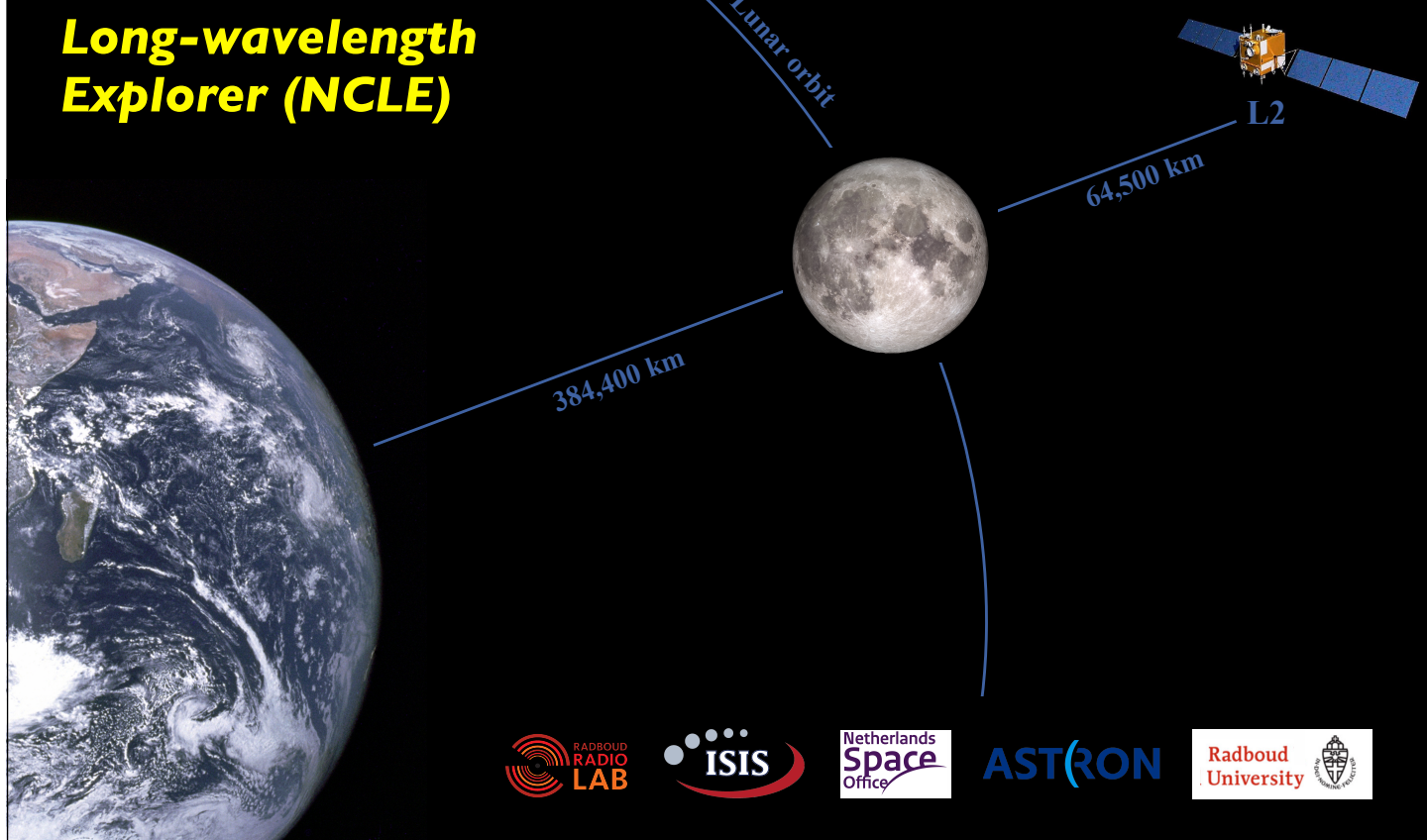


Launch in Q3 2018!

Low-frequency Radio from Space

**Netherlands – China
Long-wavelength
Explorer (NCLE)**

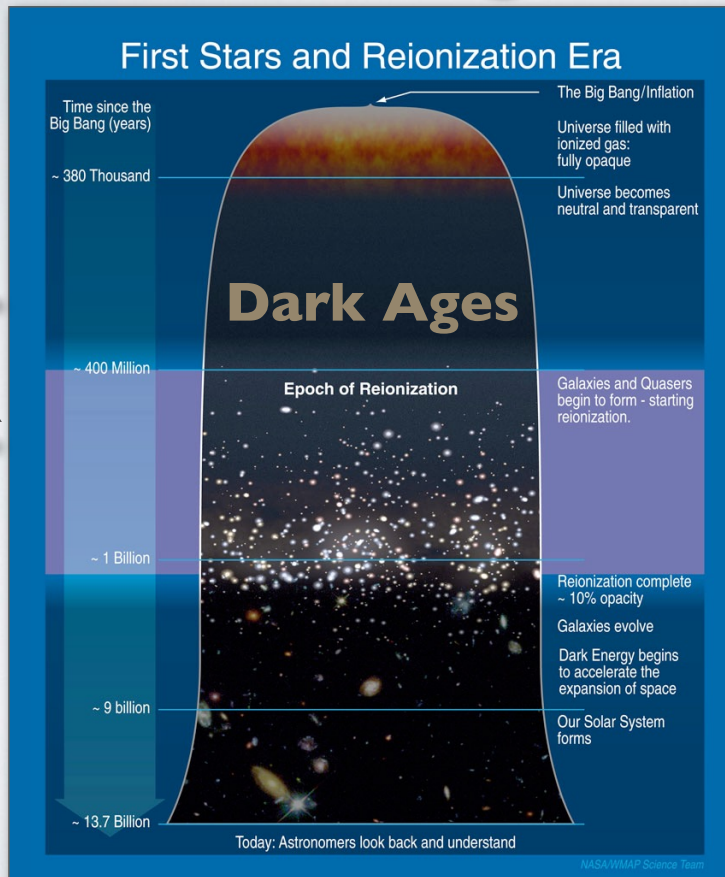
Chang'e 4 Orbiter



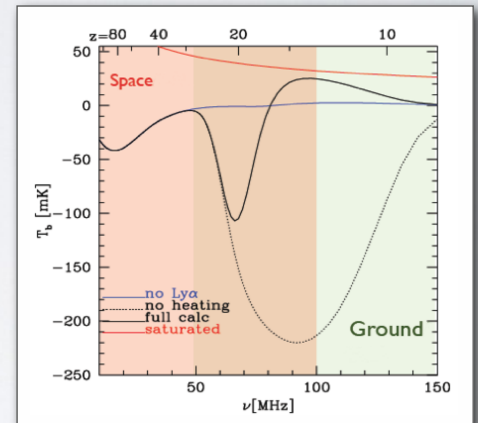
Dark Ages signal

NCLE
→

LOFAR
→



中-荷长波探测器
the Netherlands-China Long-wavelength Explorer

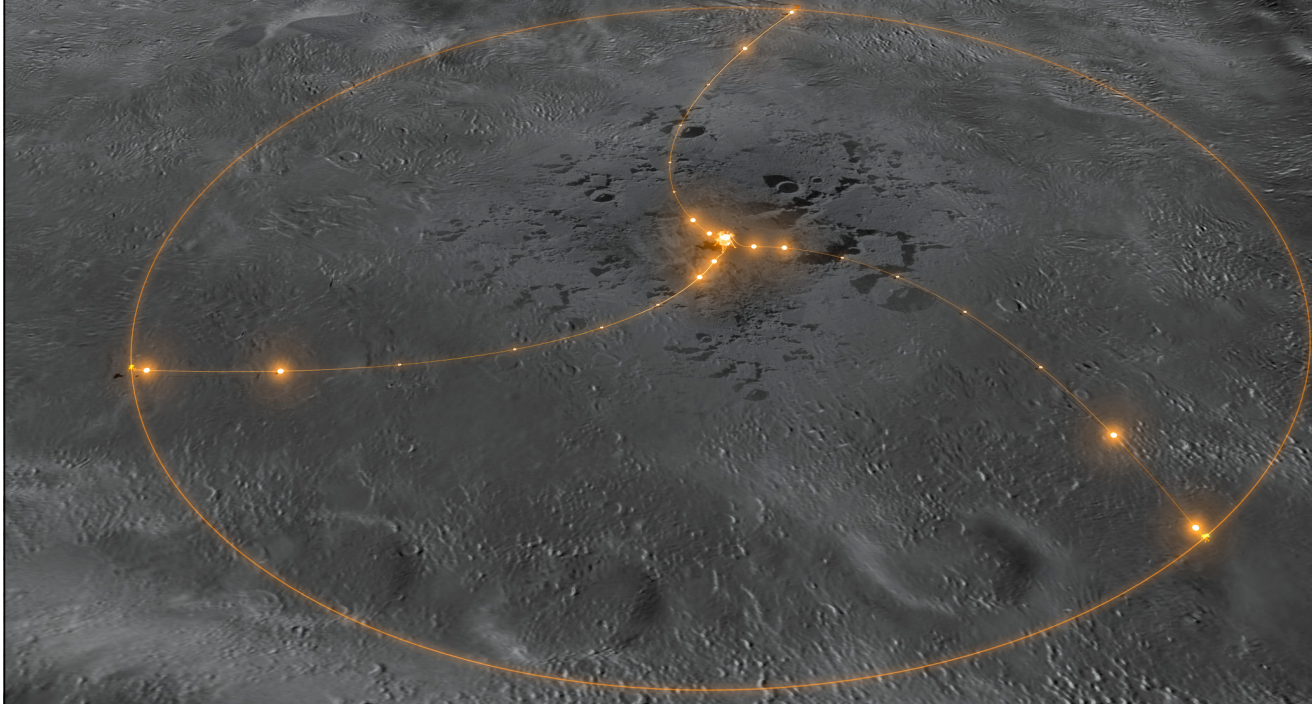


21 cm H-I line can be seen between 9 -200 MHz with one tripole and integrating for a number of months (Jester & Falcke 2009)

With the LRX we can try to detect the signal of the Dark Ages, i.e. the 21-cm line of H-I, which in the end will tell us something about the distribution of matter in the early universe. Here the LRX is a pathfinder mission: a future large lunar array will be able to detect density differences in detail. Note that this is comparable with the WMAP missions to detect the cosmic background radiation: first step was to detect it, the next to to detail observations. The LRX will be a first step, but will provide a unique image of the early universe and will provide clues on star and galaxy formation

Dark ages is period between epoch of recombination when the universe became neutral and opaque, and the epoch of reionization when the universe re-ionized due to the first stars. The HI 21 cm line is the only source of light in the Dark Ages, and provides a clue on the distribution of matter in the early universe right up to the period when the first stars are born. This addresses a key cosmological issue!

Lunar Radio Astronomy

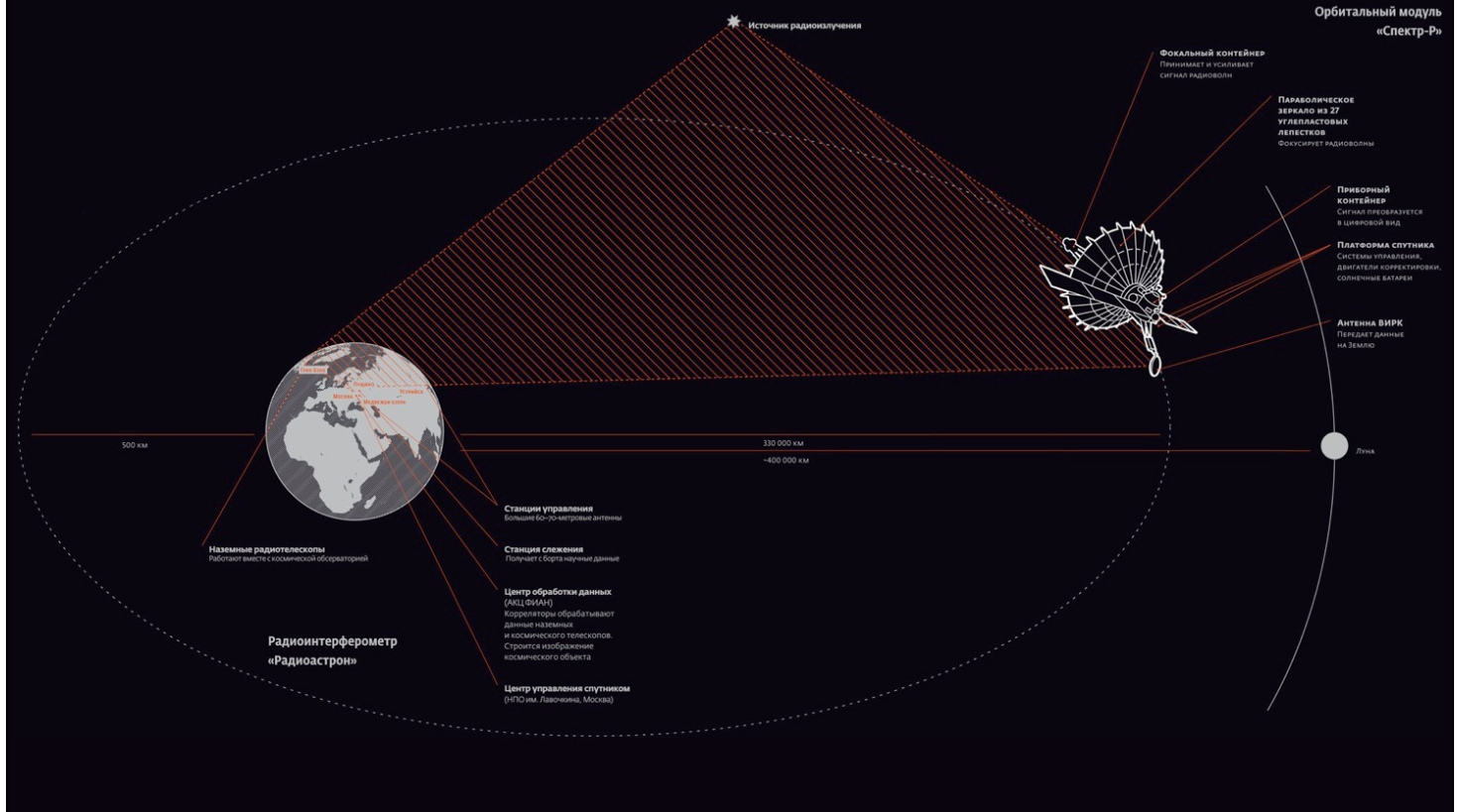


Moving beyond earth, we could imagine building an array on the moon.

By building on the dark side of the moon, we could shield the array from Earth's RFI environment.

Studies are already underway for designs for such a lunar array.

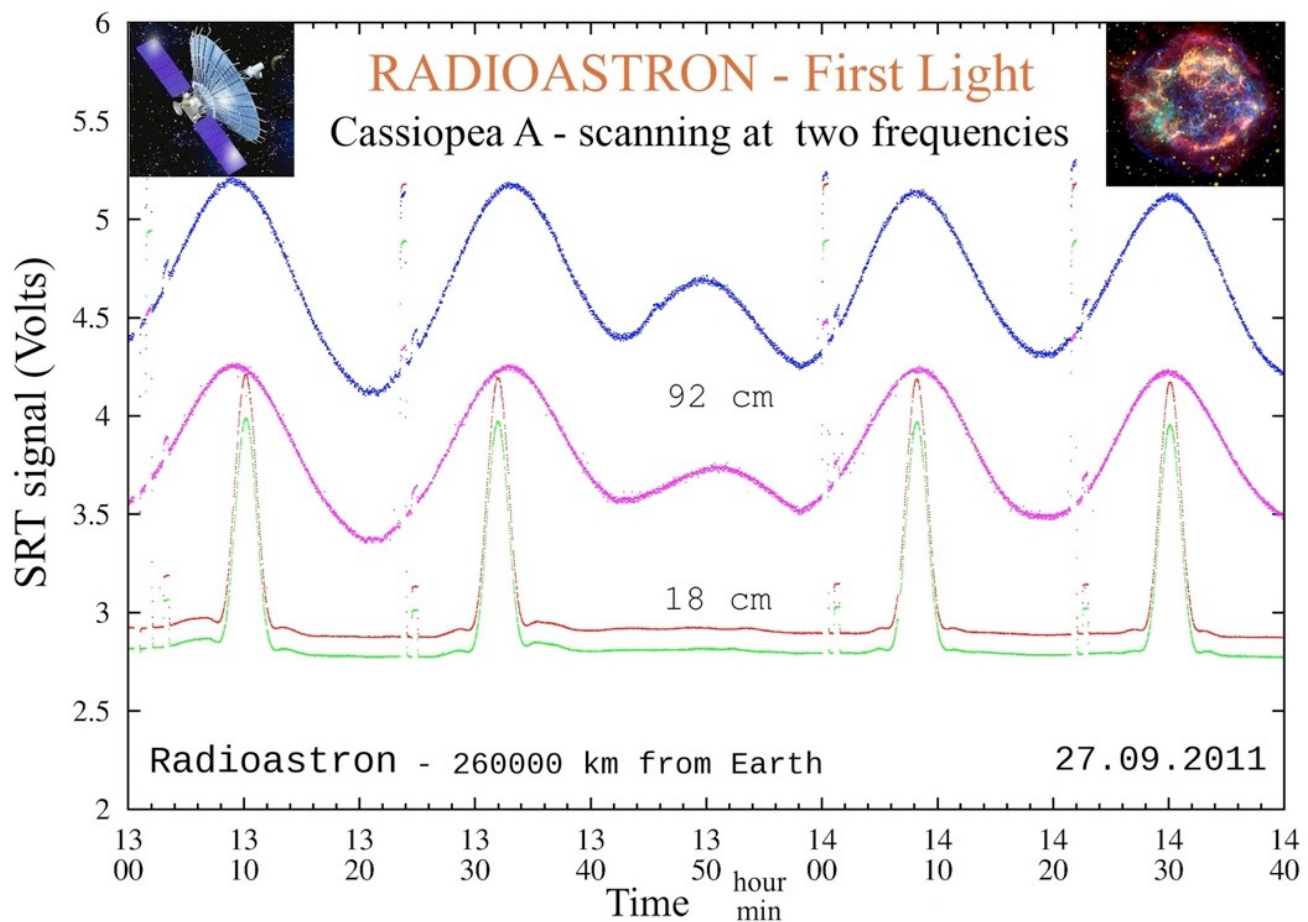
VLBI from Space



The ultimate in space-based radio astronomy would be VLBI from space.

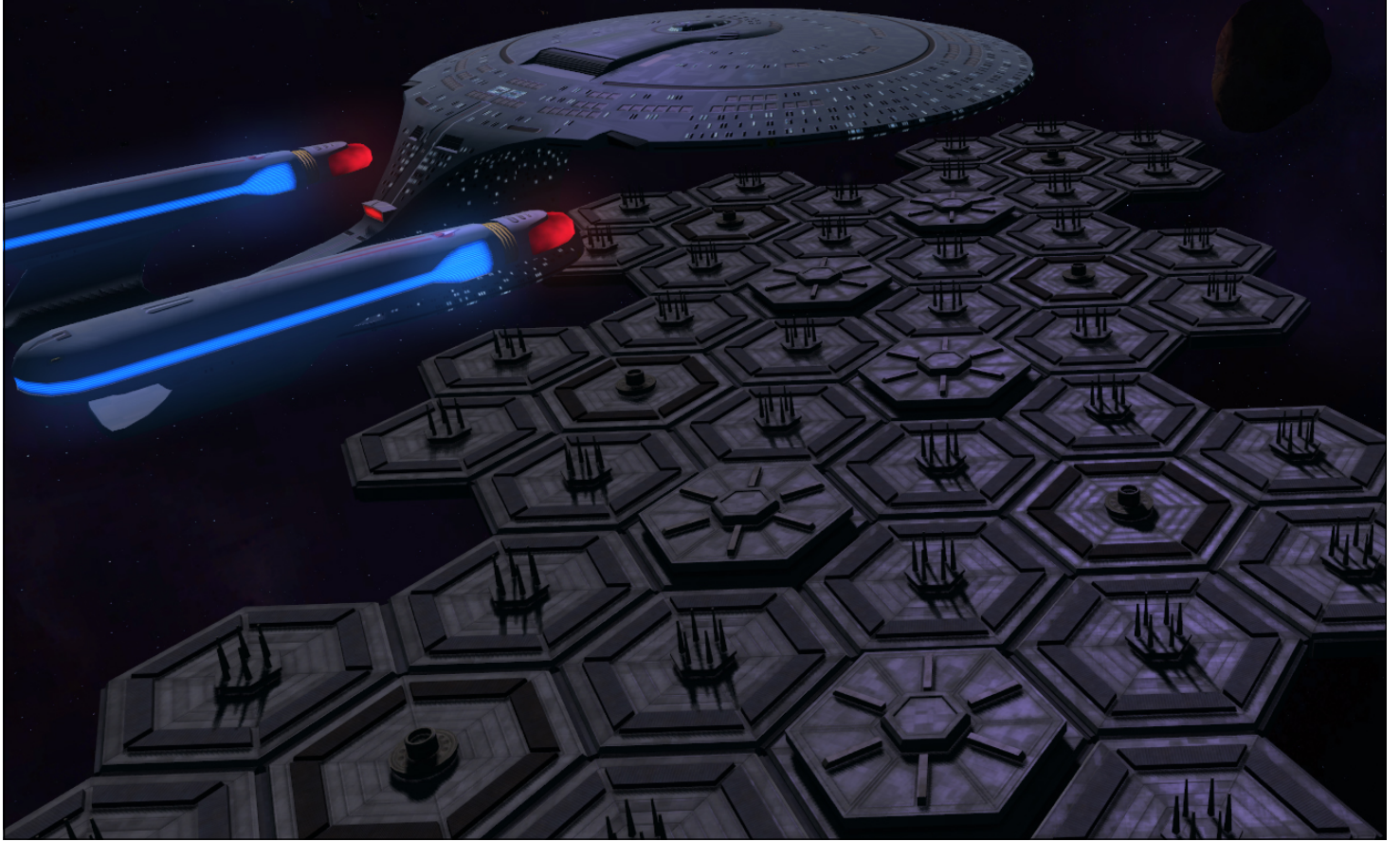
Using an array of radio satellites distributed throughout the solar system we could achieve ultra-high resolution imaging.

This sort of very high resolution imaging is necessary for example to study the environment in the immediate vicinity of extragalactic black holes.



RADIOASTRON is a first example of VLBI from space. This plot shows fringes at two different frequencies between Earth and the RADIOASTRON satellite at a distance of over 260,000 km.

...and beyond



The Future of Radio Astronomy is YOU!



Questions?